

Hydrodynamical modelling of Type IIb SNe.[★]

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ABSTRACT

We present HYDE, a new one-dimensional hydrodynamical code, and use it to construct a grid of supernova (SN) models based on solar-metallicity bare helium-core models evolved to the verge of core-collapse with MESA STAR. This grid is suited to model Type IIb SNe, which progenitor stars are thought to have lost all but a tiny fraction of their hydrogen envelopes. As previously demonstrated, such an envelope only affects the early lightcurve, and the diffusion phase and the early tail phase lightcurves are governed by the helium core. Relatively massive hydrogen envelopes do, however, affect the photospheric velocities during the diffusion phase, which could lead to underestimates of the explosion energy. Using an automated procedure we fit the bolometric lightcurves and photospheric velocities for a large sample of (17) Type IIb SNe to the grid of SN models. We find that the distribution of initial masses for the sample can be reasonably well described by a standard Salpeter IMF, although there is an under-population in the $>25 M_{\odot}$ range. The fractions of SNe with initial masses $<15 M_{\odot}$ and $<20 M_{\odot}$ are 56 and 81 percent, respectively, suggesting either the binary channel to dominate the production of Type IIb SNe or a serious flaw in our understanding of single-star mass-loss. We find correlations between the explosion energy, initial mass and mass of ^{56}Ni ; the explosion energy increase with initial mass and the mass of ^{56}Ni increase with explosion energy. The method used allows us to determine the errors in the model parameters arising from the observed quantities and the degeneracy of the solution. We find that an error in the distance and extinction propagates mainly to the derived mass of ^{56}Ni , whereas an error in the photospheric velocity propagates mainly to the derived helium-core mass and explosion energy. Fits using the bolometric lightcurve alone are completely degenerate along the $M_{\text{ej}}^2/E=\text{const}$ curve, whereas fits using also the photospheric velocities are quite robust for well-sampled SNe. Finally, we provide a description and tests of the HYDE code, and a discussion of the limitations of the method used.

Key words. supernovae: general

1. Introduction

Type IIb supernovae (SNe) are thought to arise from stars that have lost most of their hydrogen envelopes, either through stellar winds or through Roche-lobe overflow to a binary companion. These SNe are observationally characterized by a transition from Type II (with a hydrogen signature) to Type Ib (without a hydrogen but with a helium signature). Whether binary or single progenitor systems are dominating the production of Type IIb SNe is still debated, but for SN 1993J a companion star has been detected by direct observations (Maund et al. 2004; Fox et al. 2014). Because most of the hydrogen envelope has been lost, whereas the helium core is still intact, we expect these SNe to be well approximated by the explosions of bare helium cores, except during the early cooling phase. This method has been used by Bersten et al. (2012), and allows for estimates of the helium-core mass, explosion energy and mass of ^{56}Ni , whereas the progenitor radius can not be estimated without taking the hy-

drogen envelope into account. Type IIb SNe have the unique quality to allow an estimate of the helium-core mass which, contrary to the ejecta mass, is directly linked to the initial mass of the star. Some parameter studies have been published (e.g. Lyman et al. 2014), but are all based on approximate lightcurve modelling (e.g. Arnett 1982). The aim of this paper is to use the new hydrodynamical code HYDE to construct a grid of SN models based on bare helium-core models, and use this to estimate the progenitor and SN parameters for a large sample of Type IIb SNe. The sample consists of the Type IIb SNe from the Carnegie Supernova Project (CSP) sample of stripped-envelope SNe (Stritzinger 2015), as well as most of the Type IIb SNe that have been individually studied in the literature.

Application of hydrodynamics to SNe lightcurves was introduced in the 70ths (e.g. Falk & Arnett 1977), and since then a number of codes spanning a wide range of complexity have followed. Some implements more advanced physics, as multi-dimensional (e.g. Mueller et al. 1991) and radiation (e.g. Blinnikov et al. 1998) hydrodynamics, whereas others are one-dimensional and based on the diffusion approximation (e.g. Bersten et al. 2011). The different codes all have their different applications and no code is yet capable of modelling a core-collapse (CC) SN consistently, including all the relevant physics. HYDE belongs to the latter category, and like other simplified codes it has the benefit of being fast, which is critical when building model grids covering large volumes of parameter space. The use of model grids to determine the progenitor and SN parameters has been explored before (e.g. Litvinova & Nadezhin

[★] This manuscript is based on preliminary and proprietary data from the Carnegie Supernova Project (CSP), and is intended for publication in Astronomy and Astrophysics in coordination with the papers presenting the data. Several issues also remain to be solved, including a comparison with STELLA (Blinnikov et al. 1998) to determine how the calibration of the opacity floor varies within the parameter space of the model grid. The results for SNe 2003bg, 1996cb, 2011ei, and 2011fu are still based on an earlier version of the model grid using (mass) scaled versions of the $4 M_{\odot}$ bare helium-core model from Nomoto & Hashimoto (1988)

1983, 1985) with somewhat mixed results (e.g. Hamuy 2003), but the decreasing computational cost and the increasing amount of data, motivate a renewed interest in this approach. A model grid also allows the degeneracy of the solution and the errors in the progenitor and SN parameters to be estimated.

The paper is organized as follows. In Sect. 2 we describe and test the HYDE code. In Sect. 3 we describe the grid of bare helium-core and SN models, and discuss the dependence of the observed properties on the progenitor and SN parameters. In Sect. 4 we present models with low-mass hydrogen envelopes, and discuss the effects of these on the observed properties. In Sect. 5 we describe our fitting procedure, use the grid of SN models to estimate the progenitor and SN parameters for our sample of Type IIb SNe, and discuss the total sample statistics. The observational details for the Type IIb SNe from the CSP sample and those individually studied in the literature are given in Appendices A and B, respectively. Finally, we conclude and summarize the paper in Sect. 6.

2. The HYDE code

HYDE is a 1-D (spherically symmetric) hydrodynamical code based on the diffusion approximation, developed along the lines described in Falk & Arnett (1977). The code is written in C++ and may also be run in homologous mode, where the dynamics has been switched off and the thermodynamical state is solved for given the constraint of homologous expansion. The code is configurable in a number of other ways, e.g. with respect to the use of a flux-limiter and the form of the momentum and energy equations, and atomic and opacity data are read from files in generic (but proprietary) formats, and can therefore easily be updated.

2.1. Hydrodynamics

HYDE solves the hydrodynamical conservation equations for mass, momentum and thermal energy, coupled with the diffusion approximation for the radiation field (Falk & Arnett 1977, eqs. 1-4). To limit the diffusion velocity in the optically thin regime, HYDE provides an option to use a flux limiter following the prescription by Bersten et al. (2011), which is then added in the diffusion equation (eq. 4). The flux limiter transforms the radiation field from the optically thick diffusion limit to the free-streaming unidirectional limit, but is only qualitatively correct in the intermediate region (Mihalas & Weibel Mihalas 1984), and may introduce inconsistencies in the radiation field (Epstein 1981). These inconsistencies arise in the radiation pressure terms in the momentum and thermal energy equations, as the flux-limiter enforces a transformation from an isotropic radiation field to an unidirectional one, i.e. the Eddington factor ($f_E = P_R/E_R$) increases from 1/3 to 1. Therefore, HYDE provides an option to use alternative forms of the momentum and thermal energy equations, modified to be consistent with the flux limiter. To make the momentum equation consistent, we add the isotropy related factor $(3P_R - E_R)/r$ (Mihalas & Weibel Mihalas 1984, eq. 96.3), and use the time-independent first-order moment equation to rewrite the divergence of the radiation pressure in terms of radiation flux. To make the thermal energy equation consistent, we add the isotropy related factor $(3P_R - E_R)v/(pr)$ (Mihalas & Weibel Mihalas 1984, e.q. 96.9), and use the time-independent first-order moment equation to determine the Eddington factor for a radiation field with the energy density of a blackbody and the flux given by the flux limiter. This, in turn, is used to rewrite the radiation pressure in terms of radiation energy density and

the Eddington factor, and we arrive at the following modified versions of eqs. 1-4 in Falk & Arnett (1977)

$$m = \int 4\pi r^2 \rho dr \quad (1)$$

$$\frac{\partial v}{\partial t} = -4\pi r^2 \frac{\partial P_G}{\partial m} + \frac{\rho \kappa_F}{4\pi r^2 c} L + \frac{Gm}{r^2} \quad (2)$$

$$\frac{\partial}{\partial t} \left(E_G + \frac{E_R}{\rho} \right) = -(P_G + f_E E_R) \frac{\partial}{\partial t} \left(\frac{1}{\rho} \right) + (3f_E - 1) E_R \frac{v}{\rho r} - \frac{\partial L}{\partial m} + \epsilon \quad (3)$$

$$L = -(4\pi r^2)^2 \frac{\lambda ac}{3\kappa_R} \frac{\partial T^4}{\partial m} \quad (4)$$

$$\frac{\partial (f_E E_R)}{\partial r} + (3f_E - 1) \frac{E_R}{r} = -\frac{\rho \kappa_E}{4\pi r^2 c} L \quad (5)$$

where E_G and P_G are the specific (thermal) energy and the pressure for the gas, E_R the energy density for the radiation field, λ the flux limiter, ϵ the specific radioactive heating rate, and κ_R and κ_F the Rosseland mean opacity and the mean opacity weighted in energy flux. Note that the Eddington factor is determined by the form of the flux limiter, and may differ from what would be obtained from geometrical considerations

The equation of state (EOS) for the radiation field is assumed to be that of blackbody radiation, although the Eddington factor is allowed to vary between 1/3 and 1 as described. The EOS for the gas is assumed to be that of an ideal gas, including the effects of ionization, but excluding those of excitation, and degeneracy of the electron gas is not taken into account. The ionization fractions and the electron density required in the EOS are calculated by solving the Saha equation, using up to five ionization stages for each element (this is a configurable). To solve the equations we also need to specify initial and boundary conditions. At the inner boundary we adopt $L = 0$ and $v = 0$, and at the outer boundary $P_G = 0$ and $T^4 = (3/4)T_{\text{eff}}^4(\tau + q)$. The value of q may be set to 2/3 or 1/3, which gives the Eddington approximation or a modified version of it, chosen to be consistent with the flux limiter in the sense that the unidirectional limit ($f_E = 1$) is recovered when $\tau \rightarrow 0$. The initial conditions are determined by the stellar model, and for consistency the temperature structure of the model may be recalculated using the HYDE EOS assuming hydrostatic equilibrium. The zero-velocity boundary represents the division (mass cut) between the ejected material and the compact remnant, and the explosion energy is injected near this boundary in the form of thermal energy (thermal bomb), represented as an additional heating term in Eq. 3.

Given the initial and boundary conditions, the EOS, the opacity (Sect. 2.2) and the radioactive heating (Sect. 2.3), the equation system (Eqs. 1-4) is solved by a finite difference scheme similar to the one described by Falk & Arnett (1977, eqs. A1-A12). To handle strong velocity gradients (shocks) an artificial viscosity following the prescription by Von Neumann & Richtmyer (1950) is used, and added to the pressure terms in the momentum and thermal energy equations (Eqs. 2 and 3). The dynamical state is solved for using the momentum equation (Eq. 2) and a forward difference scheme, where the new state is

explicitly determined by the previous state, and is therefore trivial to advance. The thermodynamical state is solved for using the thermal energy equation (Eq. 3) and a backward difference scheme, where the new state is implicitly determined by the previous state. This results in a non-linear equation system, which is solved by a Newton-Raphson like method, where the equation system is linearised in terms of temperature corrections (Falk & Arnett 1977, appendix AIa). The time-step Δt for each calculation is initially set according to the Courant-Friedrich-Levy (CFL) condition $v_s \Delta t / \Delta r = 0.5$, where v_s and Δr are the sound-speed and the radial size of the cell, respectively, and is subsequently reduced if the (configurable) convergence criteria are not satisfied.

2.2. Opacity

The Rosseland mean opacity, the sole quantity that determines the coupling between the matter and the radiation field in the diffusion approximation, is interpolated from the OPAL opacity tables (Iglesias & Rogers 1996) complemented with the low temperature opacities by Alexander & Ferguson (1994). Note that in our flux-limited equations also the mean opacity weighted in energy flux appears, but we will assume the difference to be small, as is justified if the (grey) electron scattering opacity dominates. The opacity tables are calculated for a static medium in LTE, and therefore the line opacity as well as the opacity in the optically thin region may be underestimated (Sect. 3.5). To compensate for this lack of opacity, we use a minimum value for the opacity, commonly referred to as opacity floor. The value of this opacity floor is set to $0.01 \text{ cm}^2 \text{ g}^{-1}$ in the hydrogen envelope and $0.025 \text{ cm}^2 \text{ g}^{-1}$ in the helium core, following Bersten et al. (2012, private communication), who calibrated these values by comparison to the STELLA hydrodynamical code (Blinnikov et al. 1998).

2.3. Radioactive heating

The transfer of the γ -rays and positrons emitted in the decay chains of ^{56}Ni , ^{57}Ni and ^{44}Ti is calculated with a Monte-Carlo method similar to that by Jerkstrand et al. (2011, 2012), and the mass fractions of the isotopes evolved at each time step. The grey opacities, luminosities and decay times used are the same as in Jerkstrand et al. (2011, 2012). The deposited decay luminosity is assumed to contribute only to the heating of the gas, which is a fair approximation in the optically thick region where the degree of ionization is high (Kozma & Fransson 1992), and the heating rate is fed into the energy equation.

2.4. Observed luminosity

The observed luminosity may be taken as the comoving frame luminosity at the outer boundary, but as this boundary could be accelerated to high speeds, a transformation to the observer frame may be necessary, and the light-travel time may need to be taken into account. Therefore, HYDE provides an option to calculate the observed luminosity as $L_{\text{obs}}(t_{\text{obs}}) = [1 + 2\beta(t)]L(t)$, where $\beta(t) = v(t)/c$ and $t_{\text{obs}} = t - v(t)/c$. This expression assumes a free-streaming unidirectional radiation field, and can be derived from first principles (see e.g. 99.39 in Mihalas & Weibel Mihalas 1984). HYDE also provides an option to proceed from the luminosity at some inner surface, outside which the optical depth and the radioactive energy deposition is negligible. This reduces the influence of the flux limiter, and in this case the dif-

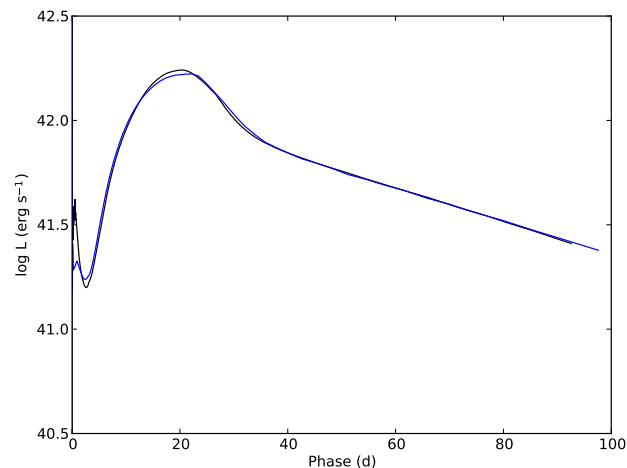


Fig. 1. Bolometric lightcurve for the $4 M_{\odot}$ bare helium-core model from Nomoto & Hashimoto (1988) as modelled with HYDE (black) and the adjusted version of the Bersten et al. (2012) He4 model presented in E14a (blue).

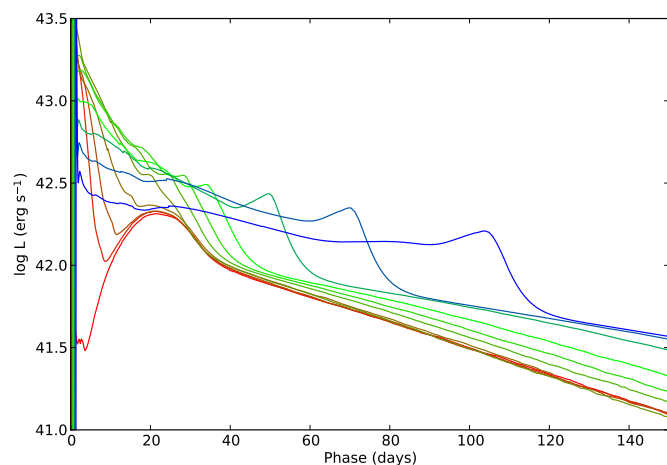


Fig. 2. Progression of model bolometric lightcurves calculated with HYDE for a $15 M_{\odot}$ MESA model with the mass-loss adjusted to yield a final mass of $11.0, 8.0, 6.0, 5.0, 4.8, 4.6, 4.4, 4.2, 4.1, 4.05$ and $4.0 M_{\odot}$, colour coded from blue ($11.0 M_{\odot}$) to red ($4.0 M_{\odot}$). The explosion parameters were $E=1.0 \times 10^{51} \text{ erg}$, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=M_{\text{He}}/M$ (Sect. 3).

ference between the luminosity in the comoving frame and that measured by a distant observer is less critical.

2.5. Tests of the code

The homologous behaviour has been tested by comparison to analytical solutions by Imshennik & Popov (1992), and the deposition of radioactive decay energy by comparison to the steady-state NLTE code described in Jerkstrand et al. (2011, 2012). Energy conservation has been tested, and is accurate to a few percent of the explosion energy in a typical run. This is illustrated by Fig. 3, which shows the change in the total energy minus the net energy gained (sum of explosion energy, radioactive heating and radiative losses), for a model with $M_{\text{He}}=4.0 M_{\odot}$, $E=1.0 \times 10^{51} \text{ erg}$, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=1.0$. This quantity, which should be zero if energy is conserved, is $<0.023 \times 10^{51} \text{ erg}$,

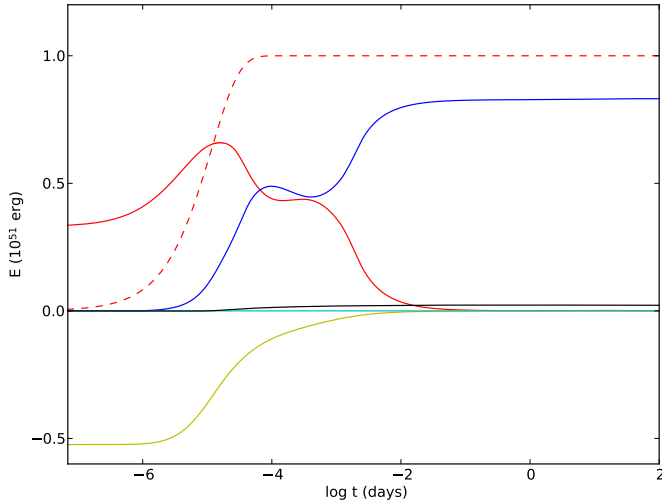


Fig. 3. The change in the total energy minus the net energy gained (black solid line) calculated with HYDE for a model with $M_{\text{He}}=4.0 M_{\odot}$, $E=1.0 \times 10^{51}$ erg, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=1.0$, where the net energy gained is given by the sum of the explosion energy, radioactive heating and radiative losses. For comparison we also show the thermal (red solid line), kinetic (blue solid line), gravitational (yellow solid line), ionization (cyan solid line) and injected explosion energy (red dashed line) for this model.

and for comparison we also show the thermal, kinetic, gravitational, ionization and injected explosion energy for this model. The hydrodynamical behaviour has been tested by comparison to the results presented in Bersten et al. (2012) using the same $4 M_{\odot}$ bare helium-core model from Nomoto & Hashimoto (1988). Figure 1 shows a comparison between the lightcurve calculated with HYDE and the lightcurve for the adjusted version of the Bersten et al. (2012) He4 model presented in Ergon et al. (2014b, hereafter E14a). Both models have the same explosion parameters ($E=1.0 \times 10^{51}$ erg, $M_{\text{Ni}}=0.075 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=0.95$), and for consistency flux-limited diffusion with $f_E = 1/3$ was used, and only ionization of hydrogen and helium was included in the EOS. Except at $\lesssim 1$ day the differences are small, but the luminosity during the first day is considerably higher. The reason for this could be differences in the zoning or the density profile, and as we have scanned the model from Nomoto & Hashimoto (1988) such differences are expected. Figure 2 shows lightcurves calculated with HYDE for a series of $15 M_{\odot}$ MESA models where the mass-loss was adjusted to yield final masses in the range $11-4 M_{\odot}$. The sequence of lightcurves shows the expected transformation from an explosion energy powered Type IIP like lightcurve to a radioactively powered Type Ib like lightcurve. Further justification for a healthy behaviour is provided in Sects. 3 and 4, where we discuss the observed and physical properties of our bare helium-core and extended models, respectively.

3. The model grid

3.1. Progenitor models

The progenitor models were constructed using MESA STAR (Paxton et al. 2011, 2013) by evolving solar-metallicity helium cores until the verge of core-collapse. This is similar to what was done in Nomoto & Hashimoto (1988), and relies on the assumption that the hydrogen envelope does not appreciably affect the evolution of the helium core. Evolving a set of solar-

Table 1. The helium, carbon-oxygen and iron core masses, the radii and the total (gravitational plus thermal) energy of the progenitor models.

M_{He} (M_{\odot})	$M_{\text{C/O}}$ (M_{\odot})	M_{Fe} (M_{\odot})	R (R_{\odot})	E_{\star} (10^{51} erg)
4.00	2.04	1.32	4.81	-0.19
4.25	2.21	1.46	4.21	-0.19
4.50	2.39	1.57	3.51	-0.30
4.75	2.58	1.61	4.22	-0.37
5.00	2.75	1.56	4.19	-0.30
5.50	3.09	1.66	3.75	-0.47
6.00	3.54	1.69	3.97	-0.49
6.50	3.92	1.77	4.13	-0.65
7.00	4.32	1.51	3.96	-0.47

metallicity $15 M_{\odot}$ models, adjusting the mass loss to yield final masses in the range $15-4 M_{\odot}$, resulted in helium cores of similar size and composition, in support of this assumption. All models are non-rotating and the Schwarzschild criteria is used for convection. Otherwise, the default MESA configuration is used, and the evolution was terminated at a central density of $10^{9.5} \text{ g cm}^{-3}$, which typically occurred slightly before core-collapse. The evolved models spans $M_{\text{He}}=4.0-5.0 M_{\odot}$ in $0.25 M_{\odot}$ steps and $M_{\text{He}}=5.0-10.0 M_{\odot}$ in $0.5 M_{\odot}$ steps, and in Table 1 we give the helium, carbon-oxygen and iron core masses, the radii and the total (gravitational plus thermal) energy for these models. Below $4.0 M_{\odot}$ the late burning stages ignited off centre, which caused convergence problems, and these stellar models were constructed by scaling of the $4.0 M_{\odot}$ density profile.

3.2. SN models

As most codes, HYDE does not include a treatment of the physics of the core-collapse itself. Instead the outcome of this event is simulated by the injection of thermal energy (thermal bomb) at some location assumed to correspond to the division between the collapsing core and the ejected material. This location is fixed to $1.5 M_{\odot}$ in all our models, and the explosion energy (E) is treated as a free parameter. HYDE does not include a network of nuclear reactions, so the explosive nuclear burning in the iron core and the inner parts of the oxygen zones, synthesizing the radioactive isotopes powering the lightcurve, can not be modelled. Because of this, and the absence of multi-dimensional effects as macroscopic mixing in 1-D (spherically symmetric) modelling, the mass (M_{Ni}) and mixing (Mix_{Ni}) of the ^{56}Ni are also treated as free parameters. The mass fraction of ^{56}Ni (X_{Ni}) was assumed to be a linearly declining function of the ejecta mass (m_{ej}) becoming zero at some fraction (Mix_{Ni}) of the total ejecta mass, expressed as $X_{\text{Ni}} \propto 1 - m_{\text{ej}}/(\text{Mix}_{\text{Ni}} M_{\text{ej}})$, $X_{\text{Ni}} \geq 0$. Note that this expression allows $\text{Mix}_{\text{Ni}} > 1$, although the interpretation of the parameter then becomes less clear. The SN explosion is thus parametrized using three parameters (E , M_{Ni} and Mix_{Ni}) and the progenitor star using one (M_{He}). The total parameter space spanned is $M_{\text{He}}=2.5-10 M_{\odot}$, $E=0.4-6.0 \times 10^{51}$ erg, $M_{\text{Ni}}=0.015-0.3 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=0.5-1.4$ using a $21 \times 24 \times 15 \times 9$ grid¹. We find this resolution to be sufficient to safely interpo-

¹ The $M_{\text{He}} > 7 M_{\odot}$ and $E > 2.2 \times 10^{51}$ erg models are not yet finalized and for SN 1996cb, 2003bg, 2011ei and 2011fu we use an older version of

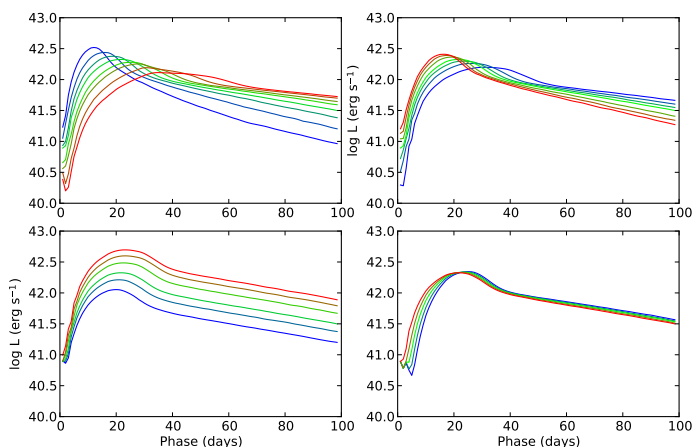


Fig. 4. Model bolometric lightcurves for day 1-100 showing the dependence on M_{He} (2.5-7.0 M_{\odot} ; upper left panel), E (0.4 - 2.2×10^{51} erg; upper right panel), M_{Ni} (0.05-0.25 M_{\odot} ; lower left panel) and Mix_{Ni} (0.6-1.0; lower right panel). Low to high values are displayed in blue to red colour coding and the values for the parameters not varied are $M_{\text{He}}=4.0 M_{\odot}$, $E=1.0 \times 10^{51}$ erg, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=1.0$.

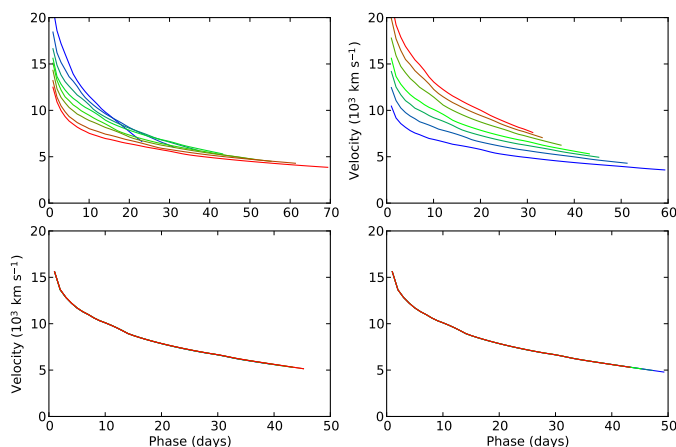


Fig. 5. Model photospheric velocities for day 1-100 showing the dependence on M_{He} (2.5-7.0 M_{\odot} ; upper left panel), E (0.4 - 2.2×10^{51} erg; upper right panel), M_{Ni} (0.05-0.25 M_{\odot} ; lower left panel) and Mix_{Ni} (0.6-1.0; lower right panel). Low to high values are displayed in blue to red colour coding and the values for the parameters not varied are $M_{\text{He}}=4.0 M_{\odot}$, $E=1.0 \times 10^{51}$ erg, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=1.0$.

late intermediate values. HYDE was configured to run with the flux-limiter and the modified momentum equation, but without calculation of the Eddington factors. The modified Eddington approximation was used at the outer boundary, and the ionization energy was not included in the EOS. The luminosity was taken from an inner surface, outside which the optical depth and the fractional radioactive energy deposition were <1 percent, and the transformation to the observer frame was ignored².

3.3. Dependence on progenitor and SN parameters

Figure 4 shows the dependence of the bolometric lightcurve on M_{He} , E , M_{Ni} and Mix_{Ni} , varying the parameters for a reference

the grid based on (mass) scaled versions of the 4.0 M_{\odot} bare helium-core model from Nomoto & Hashimoto (1988)

² Some of these settings will be changed for the final version of the model grid, but this should not have any significant effect on the results.

model with $M_{\text{He}}=4.0 M_{\odot}$, $E=1.0 \times 10^{51}$ erg, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=1.0$. Qualitatively, we expect either an increase of the explosion energy or a decrease of the ejecta mass to decrease the diffusion time for thermal radiation, to decrease the optical depth for the γ -rays emitted in the decay chain of ^{56}Ni , and to increase the expansion velocities. We therefore expect such a change to decrease the time at which peak luminosity occurs, to decrease the luminosity on the tail and to decrease the photospheric velocity. Qualitatively, we also expect the luminosity to scale with the mass of ^{56}Ni . As seen in Fig. 4, all these qualitative dependencies are well followed by the models. Quantifying the dependencies by measuring the time (t_m) and photospheric velocity (v_m) at maximum luminosity (L_m) for $\text{Mix}_{\text{Ni}}=1.0$, and fitting a power-law expression to the model grid, we get

$$\log t_m = 1.13 - 0.35 \log E_{\text{ej}} + 0.58 \log M_{\text{ej}} + 0.08 \log M_{\text{Ni}} \quad (6)$$

$$\log v_m = 1.09 + 0.43 \log E_{\text{ej}} - 0.16 \log M_{\text{ej}} \quad (7)$$

$$\log L_m = 1.34 + 0.20 \log E_{\text{ej}} - 1.02 \log M_{\text{ej}} + 0.88 \log M_{\text{Ni}} \quad (8)$$

which gives the (average) dependence of the observed quantities on the progenitor and SN parameters, expressed here for comparison in terms of the mass (M_{ej}) and energy (E_{ej}) of the ejecta. These are related to the helium core mass and explosion energy as $M_{\text{ej}} = M_{\text{He}} - M_{\text{R}}$, where M_{R} is the mass of the compact remnant (1.5 M_{\odot} ; see Sect. 3.2), and $E_{\text{ej}} = E + E_{\star}$, where E_{\star} is the total (gravitational plus thermal) energy of the progenitor model (see Table 1). Fitting the inverse relations³, we get

$$\log E_{\text{ej}} = -3.95 + 0.75 \log t_m - 0.07 \log L_m + 2.90 \log v_m \quad (9)$$

$$\log M_{\text{ej}} = -3.42 + 1.81 \log t_m - 0.18 \log L_m + 1.47 \log v_m \quad (10)$$

$$\log M_{\text{Ni}} = -4.96 + 2.08 \log t_m + 0.93 \log L_m + 1.19 \log v_m \quad (11)$$

which gives the (average) dependence of the SN and progenitor parameters on the observed quantities. Using the observed values for SN 2011dh we get values for E_{ej} , M_{ej} and M_{Ni} within ~ 30 percent of those derived by the fitting procedure in Sect. 5, and a clever parametrization of the model grid could actually be an alternative to this fitting procedure. However, given the dubious results obtained from the model grid fits for Type IIP SNe by Litvinova & Nadezhin (1983, 1985) in e.g. Hamuy (2003), care has to be taken, and we do not investigate this approach further in this work. The approximate model by Arnett (1982) is often used to infer the SN and progenitor parameters for stripped envelope SNe (e.g. Lyman et al. 2014). In this model the diffusion time and the expansion velocity depend on the mass and energy of the ejecta as $t_d \propto (M_{\text{ej}}^3/E_{\text{ej}})^{1/4}$ and $v \propto (E_{\text{ej}}/M_{\text{ej}})^{1/2}$, and inverting these gives $E_{\text{ej}} \propto t_d^2 v^3$ and $M_{\text{ej}} \propto t_d^2 v$. Comparing to the model grid fits we see that these scalings are qualitatively

³ Note that the observed quantities are not necessarily independent.

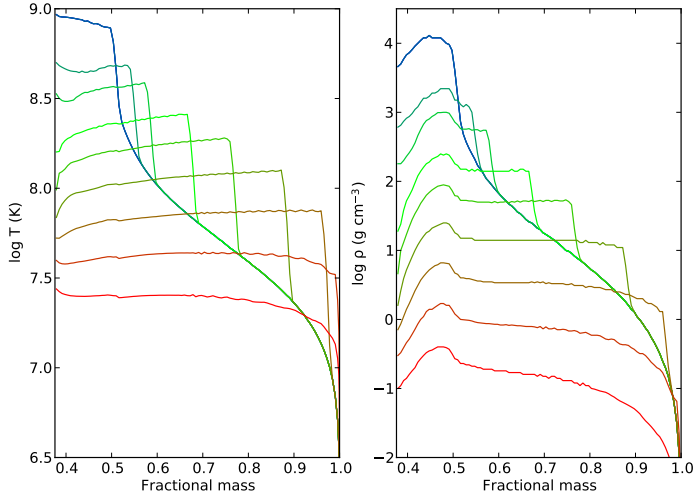


Fig. 6. Evolution of the temperature (left panel) and density (right panel) profiles between 1 and 282 seconds (shock breakout) in 10 logarithmically spaced intervals for the $4 M_{\odot}$ helium-core model, where the time has been colour coded from blue (early) to red (late).

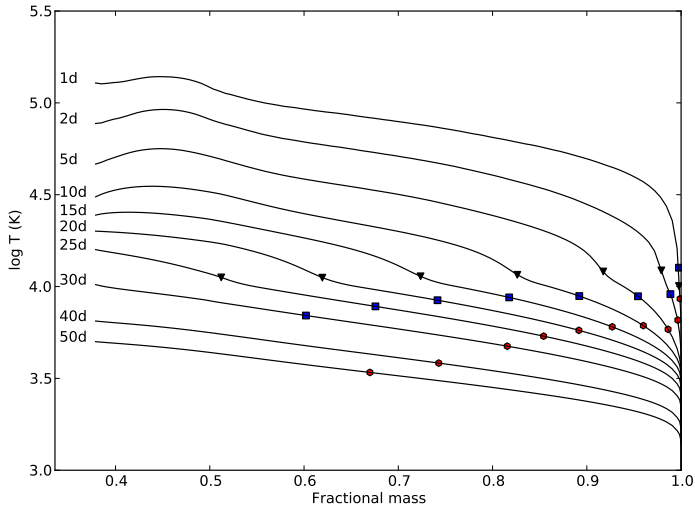


Fig. 7. Evolution of temperature profile for the $4 M_{\odot}$ helium-core model. The position of the recombination front of helium (black triangles), the photosphere (red circles) and the thermalization surface (blue squares) have been marked, and each temperature profile annotated with the time since explosion.

followed, but significant quantitative differences exist, e.g. E_{ej} is considerably less sensitive to t_{m} . Clearly, t_{m} and t_{d} and in particular v and v_{m} are not identical, but most important is likely the fact that the Arnett (1982) model assumes a constant opacity, whereas in the hydrodynamical models the (average) opacity is decreasing with time as the helium recombination front recedes through the ejecta (Sect. 3.4). An important consequence of the differences in the scalings is that t_{m} depends (roughly) on the quantity $M_{\text{ej}}^2/E_{\text{ej}}$, whereas t_{d} depends on of the quantity $M_{\text{ej}}^3/E_{\text{ej}}$ in the Arnett (1982) model. This has implications for the degeneracy of the solution, as we will discuss further in Sect. 5.5.

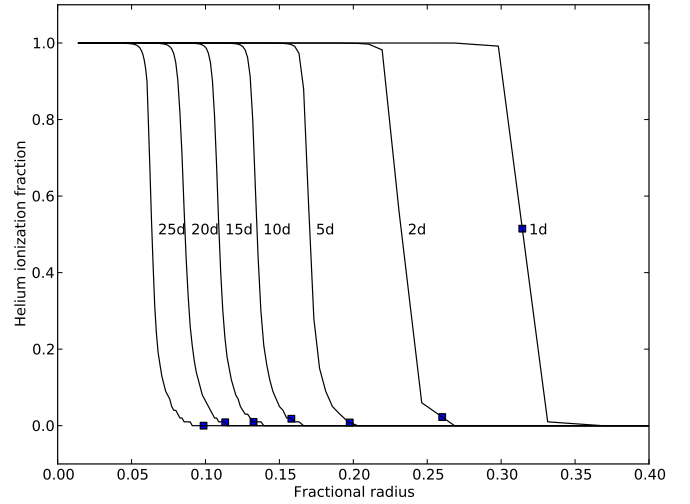


Fig. 8. Evolution of the helium ionization profile the $4 M_{\odot}$ helium-core model. The thermalization surface (blue squares) have been marked and each ionization profile annotated with the time since explosion.

3.4. Model physics

Here we discuss the physics of our bare helium-core models, exemplified by a $4 M_{\odot}$ model with SN parameters $E=1.0 \times 10^{51}$ erg, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=1.0$. We stress that the early evolution differs from that of an extended progenitor, which is discussed in Sect. 4.2.

Figure 6 shows the evolution of the density and temperature profiles from the injection of explosion energy until shock breakout, which occurs at ~ 300 seconds. The shock initially accelerates to a speed of $\sim 10000 \text{ km s}^{-1}$ in the oxygen core, but decelerates to $\sim 6000 \text{ km s}^{-1}$ in the helium envelope, where the density gradient is small. In the outermost layers the density increases drastically before it levels out in the thin convective envelope, and the shock accelerates to a speed of $\sim 30000 \text{ km s}^{-1}$ at shock breakout. The thermal and kinetic energy behind the shock is close to equipartition and the temperature high enough for the equation of state to be completely radiation dominated. During the passage of the shock through the star some thermal energy is lost due to expansion, in particular during the passage through the thin envelope, and when the radiation breaks out from the shock the thermal fraction of the energy is ~ 15 percent.

In the few minutes that follows the ejecta expand and the temperature and luminosity at the photosphere decrease rapidly by diffusion and expansion cooling. At ~ 500 seconds the outermost parts become optically thin and the photosphere starts to recede into the ejecta. At ~ 3 hours helium starts to recombine and at ~ 10 hours the recombination front overruns the photosphere. Subsequently the position of the photosphere is determined by the recombination front, slowly moving inwards in mass coordinates, but outwards in radial coordinates. Figures 7 and 8 show the evolution of the temperature and the helium ionization profile between 1 and 50 and 1 and 25 days, respectively, where we have also marked the positions of the photosphere, the thermalization surface and the recombination front. The thermalization surface, here defined as $\sqrt{3\tau_{\text{abs}}\tau_{\text{tot}}} = 2/3$ (Ensmann & Burrows 1992), is located near the outer edge of the recombination front, and follows the evolution of this until ~ 25 days when the helium has recombined. During this period the temperature at the thermalization surface is roughly constant, and declines only slowly

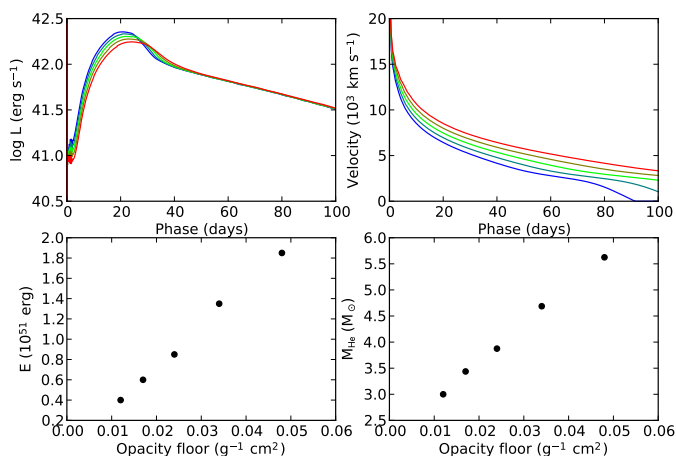


Fig. 9. Upper panels: Model bolometric lightcurves (left panel) and photospheric velocities (right panel) for day 1-100 showing the dependence on the opacity floor (0.012 - $0.048 \text{ g}^{-1} \text{ cm}^2$). Low to high values are displayed in blue to red colour coding and the model parameters are $M_{\text{He}}=4.0 M_{\odot}$, $E=1.0 \times 10^{51} \text{ erg}$, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=1.0$. Lower panels: Sensitivity of the explosion energy (left panel) and helium-core mass (right panel) to the opacity floor, calculated by fitting the model lightcurves and photospheric velocities to the model grid using the procedure described in Sect. 5.2.

from $\sim 9000 \text{ K}$ to $\sim 8000 \text{ K}$, a few thousand degrees below the temperature at the centre of the recombination front.

We note that this temperature is in good agreement with the blackbody temperature measured for SN 2011dh in E14a. Piro & Morozova (2014) argue that the blackbody temperature for SN 2011dh and other similar SNe is too low to ionize helium and that their helium envelopes might be effectively transparent, but according to our results this is not the case. Instead, we find that the gradual recombination of the helium envelope is actually what shapes the major part of the diffusion peak lightcurve for a Type IIb SN.

3.5. Model limitations

Although more sophisticated than the Arnett (1982) like models often used in SN sample studies, the hydrodynamical models used in this study still suffer from a number of limitations. Below we discuss briefly the most important of these and the limitations in our treatment of the opacity in some more detail, as the bolometric lightcurves depend critically on this quantity.

Progenitor models The progenitor models only differ in helium-core mass, and are all non-rotating and with solar metallicity. This is clearly a simplification, but as the number of SN models would increase drastically we have chosen not to vary these progenitor parameters. The effect of a low-mass hydrogen-envelope, not present in our bare helium-core models, is discussed separately in Sect. 4.

Hydrodynamics As HYDE does not include a nuclear reaction network the effect of the explosive nucleosynthesis in the inner part of the ejecta is not included, and as HYDE is 1-D, macroscopic mixing of the nuclear burning zones (see Iwamoto et al. 1997) is prohibited. The mass-cut is fixed at $1.5 M_{\odot}$ and although fallback of material onto the compact remnant is not prohibited, the artificial zero-velocity inner boundary condition

will cause this material to bounce. To properly handle fallback a piston-driven explosion without an inner boundary condition on the velocity would be needed.

Radiative transfer The optically thin regime is not handled correctly as the code is based on the flux-limited diffusion approximation (e.g. Epstein 1981). The treatment of the optically thin region is critical when calculating spectra or broad-band photometry, but probably of less importance when calculating the bolometric lightcurve. A correct treatment of the optically thin region could also be important for the radiative acceleration of the outer parts of the ejecta occurring at early times. This could effect the bolometric lightcurve during the cooling phase, but probably not later on.

Opacity One major limitation with HYDE is the absence of a proper treatment of the line (bound-bound) opacity. The opacity tables used are calculated for a static medium and does not take into account the effect of a velocity gradient, which tend to increase the line opacity (e.g. Karp et al. 1977). Furthermore, as the opacity is calculate for a medium in LTE, it may not apply in the optically thin region, where non-thermal ionization could increase the electron scattering contribution. To compensate for this lack of opacity, HYDE makes use of an opacity floor (Sect. 2.2). The value of this floor is set to $0.024 \text{ g}^{-1} \text{ cm}^2$, which is much lower than the electron scattering opacity of $\sim 0.2 \text{ g}^{-1} \text{ cm}^2$ for fully ionized helium-core material, and only affects the region outside the recombination front (Sect. 3.4). The upper panels of Fig. 9 show the dependence of the model lightcurves and photospheric velocities on the opacity floor for a $4 M_{\odot}$ model with explosion parameters $E=1.0 \times 10^{51} \text{ erg}$, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=1.0$. These are not particularly sensitive, and a doubling of the opacity floor corresponds to an increase in the photospheric velocities and a shift to later times of the diffusion peak of 10-15 percent. The lower panels of Fig. 9 show the sensitivity of the estimated helium-core mass and explosion energy to the opacity floor, calculated by fitting the model lightcurves and photospheric velocities to the model grid using the same procedure as in Sect. 5. The explosion energy depends strongly on the photospheric velocities (Eqs. 9), so this quantity is rather sensitive and almost proportional to the opacity floor, whereas the helium-core mass is less affected and a doubling of the opacity floor corresponds to an increase of ~ 25 percent. A proper investigation of the effects of the opacity floor on our results in Sect. 5 would require a thorough comparison with a code capable of calculating the line opacity correctly, and is outside the scope of this paper. We conclude, however, that the helium-core mass is not particularly sensitive to the choice of opacity floor, whereas this choice is more critical with respect to the explosion energy.

4. The hydrogen envelope

As our aim is to use the grid of bare helium-core models to fit the bolometric lightcurves and the photospheric velocities of Type IIb SNe, which may have extended low-mass hydrogen envelopes surrounding their helium cores, it is of importance to investigate which effect such envelopes would have on the observed properties, as well as on the results obtained in Sect. 5. It is also of interest to discuss the physics of such models and compare to the physics of bare helium (Sect. 3.4).

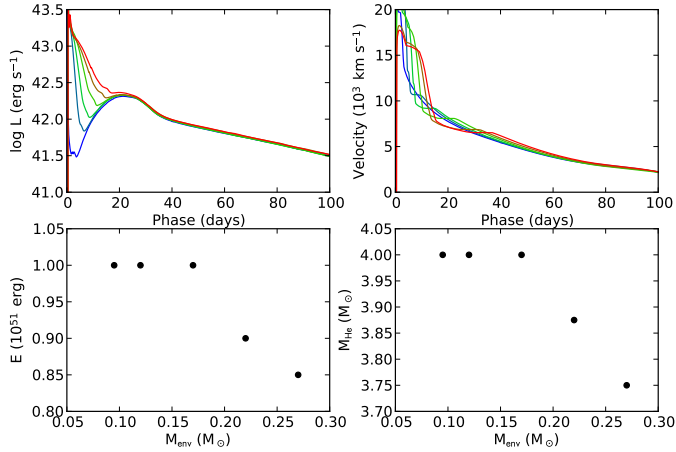


Fig. 10. Upper panels: Progression of model bolometric lightcurves (left panel) and photospheric velocities (right panel) calculated with HYDE for a $15 M_{\odot}$ MESA model with the mass-loss adjusted to yield a final mass of 4.2, 4.15, 4.1, 4.05, 4.025 and $4.0 M_{\odot}$, colour coded from red ($4.2 M_{\odot}$) to blue ($4.0 M_{\odot}$). The SN parameters were $E=1.0 \times 10^{51}$ erg, $M_{\text{Ni}}=0.1 M_{\odot}$ and $\text{Mix}_{\text{Ni}}=M_{\text{He}}/M$. Lower panels: Sensitivity of the explosion energy (left panel) and helium-core mass (right panel) to the mass of the hydrogen envelope, calculated by fitting the model lightcurves and photospheric velocities to the model grid using the procedure described in Sect. 5.2.

4.1. Effect on the observed properties

The upper panels of Fig. 10 show the bolometric lightcurve and photospheric velocities for a sequence of $15 M_{\odot}$ MESA models, where the mass loss was adjusted to yield final masses in the range $4.2\text{--}4.0 M_{\odot}$. Defining the hydrogen envelope to begin where $X > 0.01$, this corresponds to hydrogen envelope masses in the range $0.27\text{--}0.07 M_{\odot}$. The bolometric lightcurves for all models show an initial decline phase corresponding to the cooling of the thermal explosion energy deposited in the hydrogen envelope, the length of which decreases with decreasing mass of the envelope. The reason for this is twofold, first the thermal energy deposited in the hydrogen envelope decrease with the mass of it, and secondly the radius of the progenitor stars decrease, decreasing the time scale for expansion cooling. Models for Type IIb SNe often have an increased helium abundance in the hydrogen envelope (e.g. Woosley et al. 1994), due to mixing of helium into the base of the hydrogen envelope. This results in smaller progenitor radii due to decreased opacities, and therefore in shorter durations of the cooling phase. Our models have lower helium abundances as compared to Woosley et al. (1994) and Shigeyama et al. (1994), which should be kept in mind. During most of the cooling phase the photospheric velocities are much higher than those for a bare helium-core model, but decrease quickly at the luminosity minimum, after which follows a period when they are significantly lower. The latter effect is larger for models with more massive hydrogen envelopes and is likely caused by deceleration of the helium core. As suggested by Eqs. 9-10, and as demonstrated in Sect. 5.4, the sensitivity of the estimated explosion energy to an error in the photospheric velocity is high ($E \sim v^3$), whereas the sensitivity of the helium-core mass is lower. This is quantified by the lower panels of Fig. 10, which show the sensitivity of the estimated explosion energy and helium-core mass to the mass of the hydrogen envelope, calculated by fitting the model lightcurves and photospheric velocities to the model grid using the same procedure as in Sect. 5. If the mass of the

hydrogen envelope is larger than $\sim 0.2 M_{\odot}$ the explosion energy is significantly underestimated, whereas the helium-core mass is less affected. Therefore the presence of a relatively massive hydrogen envelope could have a significant effect on the estimated explosion energy when using bare helium-core models. It is worth noting that if we subtract the energy in the hydrogen envelope, the ejecta energy is in much better agreement with what is estimated from the fit, which therefore rather provide an estimate of the energy in the helium core. Except for the effect on the photospheric velocities, the presence of the hydrogen envelope does not seem to appreciably affect the observed properties after the luminosity minimum.

4.2. Model physics

Here we discuss the physics of models with low-mass hydrogen envelopes, exemplified by the $4.05 M_{\odot}$ model shown in Fig. 10. This model has a hydrogen envelope of $0.17 M_{\odot}$, an average hydrogen fraction in the envelope of 0.5, and reach the luminosity minimum at ~ 11 days, which is similar to, but slightly later than was observed for SN 1993J.

The passage of the shock through the helium core proceeds as described for the bare helium-core model, and in the steep density gradient between the helium core and the hydrogen envelope it accelerates to $\sim 20000 \text{ km s}^{-1}$. Once in the hydrogen envelope, where the density is roughly constant, the shock gradually decelerates to $\sim 6000 \text{ km s}^{-1}$, which gives rise to a reverse shock propagating backwards into the helium envelope. During the passage of the shock through the hydrogen envelope, the helium core expands and most of the deposited thermal energy is cooled away. At shock breakout the (relatively) cool and expanded helium core is surrounded by the hot and compressed hydrogen envelope, and the subsequent evolution is determined by the expansion and cooling of this envelope. The fraction of the energy deposited in the envelope is about 10 percent, roughly equipartitioned into thermal and kinetic energy.

Figure 11 shows the evolution of the density and temperature profiles from shock breakout, which occurs at ~ 0.3 days, until the luminosity minimum. The profiles are similar to those obtained by modelling of SN 1993J in Woosley et al. (1994), Shigeyama et al. (1994) and Blinnikov et al. (1998). Initially the hydrogen envelope is opaque and ionized, and the surface luminosity and temperature decreasing by expansion cooling, but at ~ 4 days the outer parts become optically thin and the photosphere starts to recede into the ejecta. The helium starts to recombine at about the same time, whereas hydrogen stays ionized until ~ 7 days, and at about ~ 6 days the photosphere starts to trace the helium recombination front as in the bare helium-core models. At the luminosity minimum the hydrogen in the envelope has recombined, the photosphere is located close to the interface between the hydrogen and helium envelope, and the temperature at the thermalization surface is $\sim 9500 \text{ K}$.

5. Model grid fits

Here we use an automated procedure to fit the bolometric lightcurves and photospheric velocities for the CSP and literature samples of Type IIb SNe to those of our grid of SN models. Our method allows us to determine the sensitivity of the derived quantities to errors in the observed quantities, as well as to investigate the degeneracy of the solutions found. As discussed, the grid is based on bare helium-core models, and given that a low-mass hydrogen envelope would mainly affect the cooling phase (Sect. 4.1), the diffusion phase and early tail lightcurves,

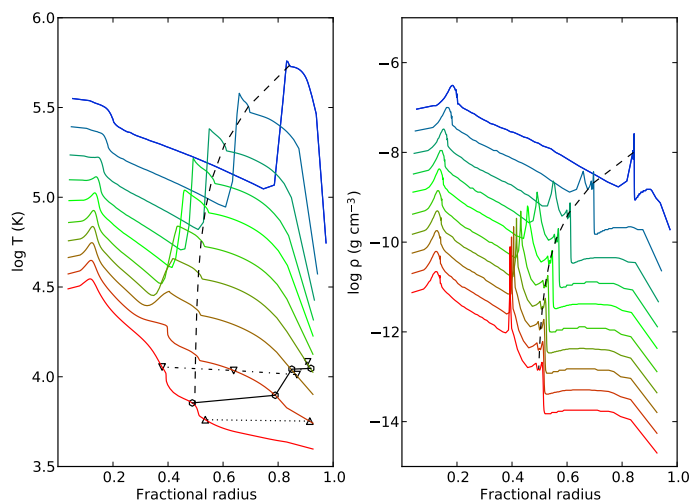


Fig. 11. Evolution of the temperature (left panel) and density (right panel) profiles between shock breakout (0.3 days) and the luminosity minimum (11 days) in 10 logarithmically spaced intervals for the $4.05 M_{\odot}$ MESA model, where the time has been colour coded from blue (early) to red (late). The interface between the helium and hydrogen envelopes (dashed line), the photosphere (circles and dashed line) and the helium (downward triangles and dot-dashed line) and hydrogen (upward triangles and dotted line) recombination fronts are also shown.

and the diffusion phase photospheric velocities, are used to determine the helium core and SN parameters. However, keep in mind that relatively massive hydrogen envelopes do affect the photospheric velocities in the diffusion phase, which might lead to significant underestimates of the explosion energy (Sect. 4.1). We also make the assumption, justified for Type IIb SNe, that the helium core is not affected by mass loss.

5.1. The Type IIb SNe sample

The Type IIb SNe in the CSP stripped-envelope SNe sample (Stritzinger 2015) consist of SNe 2004ex, 2004ff, 2005Q, 2006T, 2006ba, 2006bf, 2008aq, 2009K, 2009Z and 2009dq, of which SNe 2006bf and 2009dq have been excluded due to bad sampling of the lightcurves. The Type IIb SNe that have been individually studied in the literature consist of SNe 1993J (e.g. Richmond et al. 1994, 1996), 1996cb (Qiu et al. 1999), 2003bg (Hamuy et al. 2009), 2008ax (e.g. Taubenberger et al. 2011), 2010as (Folatelli et al. 2014), 2011dh (e.g. E14a), 2011fu (Kumar et al. 2013), 2011hs (Bufano et al. 2014), 2011ei (Milisavljevic et al. 2013) and 2013df (Van Dyk et al. 2014), of which the data for SN 2010as was not yet available when this work began, and is therefore not included. Observations of 4 additional Type IIb SNe (2001gd 2006el 2008bo and 2008cw) have been published as part of surveys, but are not included in our sample. Out of the SNe in the sample, 1993J, 2008ax and 2011dh stands out by the quality of the data as well as the hard constraints on the explosion epochs. The observational details, the constraints on the explosion epochs and the adopted distances and extinctions for the CSP sample and the sample of individually studied SNe are given in Appendices A and B, respectively, where we also describe how the photospheric velocities were estimated. The pseudo-bolometric lightcurves were calculated from the photometry using the methods described in E14a, and a UV to MIR bolometric correction (BC) determined from SN 2011dh applied. The flux falling outside this wavelength range

was not corrected for, but given the results from the steady-state NLTE modelling of SN 2011dh presented in Ergon et al. (2014a, hereafter E14b), this correction is likely to be small (<0.15 mag).

5.2. Fitting procedure

The fitting is done by minimization of the square of the relative residuals, giving equal weight to the diffusion phase lightcurve, the tail lightcurve and the diffusion phase photospheric velocity evolution. The division between the diffusion and tail phases is made roughly at the point where the decline rate of the bolometric lightcurve becomes constant. If there is any sign of a cooling phase, the beginning of the diffusion phase is set to a few days after the rise to peak begins, and otherwise to the first observation. Photospheric velocities above the interface between the helium core and the hydrogen envelope are excluded from the fit. As discussed in E14a this velocity can be estimated from the minimum velocity for the $H\alpha$ absorption minimum, but is otherwise set to 10000 km s^{-1} . As the explosion epochs in many cases are not well constrained we fit, not only the progenitor (M_{He}) and SN (E , M_{Ni} and Mix_{Ni}) parameters, but also the epoch of explosion, which is allowed to vary between the hard limits obtained from detections and non-detections. The errors in the bolometric lightcurves arising from the uncertainties in distance and extinction, and a systematic error in the photospheric velocities, assumed to be 15 percent, were propagated by standard methods.

5.3. Results and comparisons

Figures 12 and 13 show the best-fit model bolometric lightcurve and photospheric velocity evolution, compared to the observed UV to MIR pseudo-bolometric lightcurve and estimated photospheric velocity evolution, as well as contour plots of the standard deviation in the fits, normalized to that of the optimal model, projected onto the E - M_{He} plane. Tables 2 and 3 give the helium-core mass, explosion energy, mass and mixing of ^{56}Ni , and explosion epoch for the best-fit models and the corresponding errors. The fits are mostly good, and for SNe 1993J, 2004ex, 2004ff, 2006T, 2006ba, 2008aq, 2008ax, 2011dh, 2011hs and 2013df the solutions are well constrained in the the E - M_{He} plane, although SN 2011hs is not well constrained below as it lies at the border of the covered parameter space. For SNe 2003bg, 2009T, 2009K and 2011ei, the constraint from the bolometric lightcurve is weak due to the limited coverage, and the solutions are quite degenerate along the $M/E=\text{const}$ curve (see Sect. 5.5). A significant degeneracy along the $M/E=\text{const}$ curve is also seen for SNe 1996cb and 2011fu, although the lightcurve coverage for these SNe is much better, whereas for SN 2009Z, the solution is a bit degenerate along the $M^2/E=\text{const}$ curve (which would suggest a weak constraint from the velocity, see Sect. 5.5).

Several SNe (1993J, 2003bg, 2008ax, 2011dh and 2011hs) in our sample have been studied using hydrodynamical modelling in other works, and eight of the SNe are also included in the sample study by Lyman et al. (2014), based on approximate Arnett (1982) models. Figure 14 shows as comparison of the helium-core mass, explosion energy and mass of ^{56}Ni estimated in these works with our results. In cases where the results are given in terms of the mass and energy of the ejecta, a mass of $1.5 M_{\odot}$ has been assumed for the compact remnant, and the difference between the explosion and ejecta energy ignored. Turning first to the results obtained with hydrodynamical modelling, the agreement is mostly reasonable, but some clear differences exist. In particular the explosion energies estimated in our study for SNe

1993J, 2003bg and 2008ax are considerably lower than in other works. SN 1993J has been modelled by Woosley et al. (1994) and Shigeyama et al. (1994) among others, SN 2003bg by Mazzali et al. (2009) and SN 2008ax by Tsvetkov et al. (2009). In the case of SNe 1993J and 2003bg the disagreement could be caused by the presence of a relatively massive hydrogen envelope (Sect. 4), as suggested by the extent of the cooling phase in the case of SN 1993J. As discussed in Sect.4.1, the explosion energy estimated from the fit is rather an estimate of the explosion energy deposited in the helium core. According to Shigeyama et al. (1994) and Mazzali et al. (2009), this energy was $\sim 0.6 \times 10^{51}$ and $\lesssim 2.5 \times 10^{51}$ erg in their models of SNe 1993J and 2003bg, respectively, which is in better agreement with our results, and in support of the hypothesis. However, in this case of SN 2008ax, the situation could be the reverse. The absence of an extended cooling phase suggests a relatively low-mass hydrogen envelope, so the SN 1993J based model used by Tsvetkov et al. (2009) could lead to a significant overestimate of the explosion energy.

For SN 2011hs, which was modelled by (Bufano et al. 2014), we find a significantly lower helium core mass, and a significantly higher mass of ^{56}Ni . The difference in the mass of ^{56}Ni can be traced back to differences in the adopted extinction and distance (Appendix B), and with respect to the helium-core mass it is worth noting that models with $M_{\text{He}} < 3.3 M_{\odot}$ was not tested by Bufano et al. (2014), so their modelling do not provide a lower bound on this quantity. For SN 2011dh our results are in good agreement with those in Bersten et al. (2012), and also with those obtained in E14b based on the <400 days lightcurve and an extended version of the model grid using a BC determined with steady-state NLTE modelling. They are also further supported by the modelling of nebular spectra in Jerkstrand et al. (2014) and the stellar evolutionary progenitor analysis by Maund et al. (2011).

Comparing to the results from the sample study by Lyman et al. (2014) we see a reasonable agreement with respect to the helium-core masses, as well as the masses of ^{56}Ni , although these differ considerably in a few cases, which can be traced back to the adopted extinctions and distances (Appendices A and B). On the other hand, the explosion energies estimated by Lyman et al. (2014) are systematically higher as compared to our results. As the (Arnett 1982) model is considerably simpler than our models, the explanation cannot be the same as discussed above, and is rather related to the unclear relation between expansion velocity and the observed velocities in the Arnett (1982) model. This results in a considerable uncertainty, propagating mainly to the estimated explosion energy (see Sect. 5.4 with respect to the hydrodynamical models). Nevertheless, the explosion energy estimated for SN 2003bg by Lyman et al. (2014) is actually in better agreement with the hydrodynamical modelling by Mazzali et al. (2009) than our results. We speculate that the reason for this is that Lyman et al. (2014) include the high photospheric velocities in the hydrogen envelope, excluded for consistency in our bare helium-core fits (Sect. 5.2), and therefore arrives at an estimate of the total kinetic energy in the ejecta. Further comparisons with the results by Lyman et al. (2014) are made in Sect. 5.6, where we discuss the sample statistics. Finally, it is worth noting that the small ejecta mass of $0.3 M_{\odot}$ estimated for SN 2011ei from modelling of nebular spectra in Milisavljevic et al. (2013) seems to be excluded by our results. This SN is quite interesting, as it shows the most extreme values for the progenitor and SN parameters in the sample. Whereas the helium-core mass and explosion energy are the highest, the mass of ^{56}Ni is the lowest. The high helium-core mass and explo-

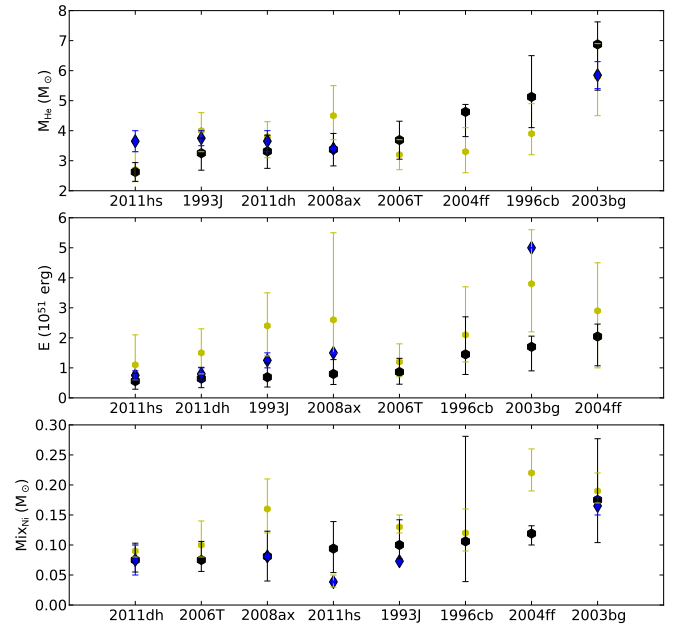


Fig. 14. Comparison of the helium-core mass (upper panel), explosion energy (middle panel) and mass of ^{56}Ni as estimated using hydrodynamical modelling in this (black circles) and other (blue diamonds) works, and as estimated using approximate Arnett (1982) models by Lyman et al. (2014) (yellow circles), where in each panel the SNe have been ordered with respect to the values estimated in this work. In cases where the results are given in terms of the mass and energy of the ejecta, a mass of $1.5 M_{\odot}$ has been assumed for the compact remnant, and the difference between the explosion and ejecta energy ignored.

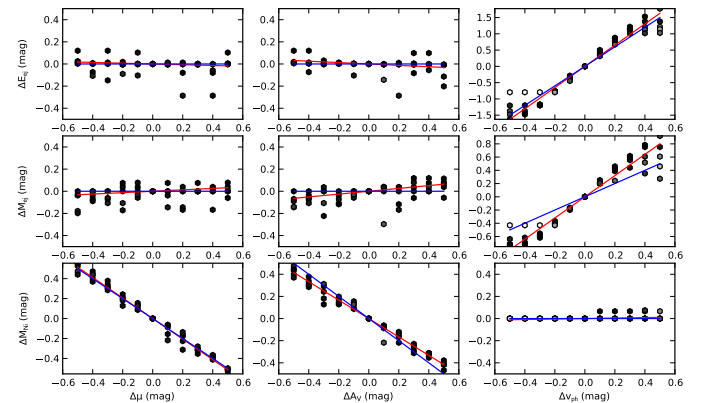


Fig. 15. Sensitivity of the derived ejecta energy (upper panels), ejecta mass (middle panels) and mass of ^{56}Ni (lower panels) to a change in the distance (left panels), extinction (middle panels) and photospheric velocities (right panels). For consistency the changes in all quantities are expressed in magnitudes. The derived quantities for our sample of Type IIb SNe are shown as black dots, power-law fits as red solid lines and the scalings expected from the Arnett (1982) model as blue solid lines.

sion energy derived stem mainly from the unusually high photospheric velocities, whereas the time at which peak luminosity occurs is quite typical.

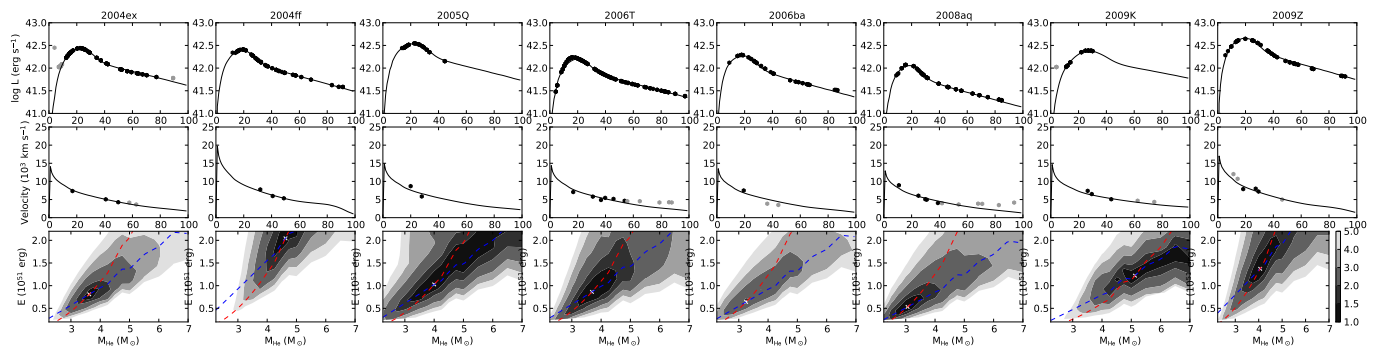


Fig. 12. Bolometric lightcurve (upper panels) and photospheric velocity evolution (middle panels) for the best-fit models as compared to the observed UV to MIR pseudo-bolometric lightcurve and estimated photospheric velocity evolution for the CSP sample of Type IIb SNe. The lower panels show contour plots of the standard deviation in the fits, normalized to that of the optimal model, projected onto the E - M_{He} plane. We also show the constraints $M_{\text{ej}}/E_{\text{ej}} = \text{const}$ (blue) and $M_{\text{ej}}^2/E_{\text{ej}} = \text{const}$ (red) provided by the photospheric velocity evolution and the bolometric lightcurve, respectively.

Table 2. Explosion energy, helium-core mass, mass and mixing of the ^{56}Ni and epoch of explosion for the best-fit models for the CSP sample of Type IIb SNe.

SN	E (10^{51} erg)	M_{He} (M_{\odot})	M_{Ni} (M_{\odot})	Mix_{Ni}	JD_{exp} (+2400000) (days)
2004ex	0.81 (+0.62,-0.36)	3.62 (+0.75,-0.62)	0.131 (+0.044,-0.031)	1.00 (+0.21,-0.10)	53285.84 (+0.50,-0.56)
2004ff	2.04 (+0.41,-0.98)	4.62 (+0.25,-0.82)	0.119 (+0.013,-0.019)	1.55 (+0.05,-0.76)	53294.89 (+1.80,-0.50)
2005Q	1.03 (+0.92,-0.38)	4.00 (+1.02,-0.62)	0.181 (+0.038,-0.026)	1.40 (+0.00,-0.49)	53385.55 (+2.92,-1.00)
2006T	0.86 (+0.46,-0.41)	3.69 (+0.63,-0.64)	0.075 (+0.031,-0.019)	1.20 (+0.40,-0.10)	53762.96 (+0.50,-1.03)
2006ba	0.64 (+0.54,-0.24)	3.25 (+0.75,-0.38)	0.088 (+0.025,-0.019)	1.60 (+0.00,-0.21)	53805.31 (+1.00,-0.50)
2008aq	0.53 (+0.22,-0.22)	3.06 (+0.25,-0.44)	0.050 (+0.025,-0.018)	1.20 (+0.00,-0.55)	54514.94 (+2.24,-0.00)
2009K	1.22 (+0.81,-0.56)	5.12 (+1.91,-1.01)	0.144 (+0.032,-0.025)	1.12 (+0.18,-0.08)	54844.08 (+0.00,-0.00)
2009Z	1.37 (+0.97,-0.57)	4.06 (+0.69,-0.62)	0.225 (+0.031,-0.032)	1.60 (+0.00,-0.05)	54859.53 (+0.50,-0.00)

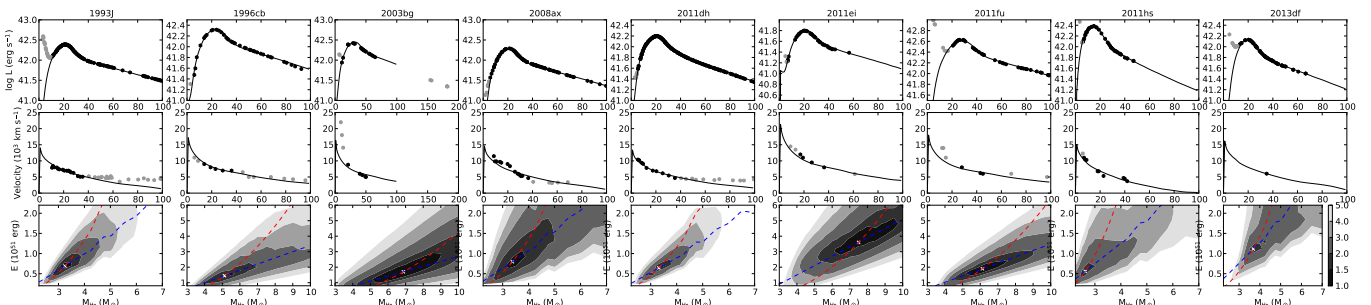


Fig. 13. Bolometric lightcurve (upper panels) and photospheric velocity evolution (middle panels) for the best-fit models as compared to the observed UV to MIR pseudo-bolometric lightcurve and estimated photospheric velocity evolution for the sample of individually studied Type IIb SNe. The lower panels show contour plots of the standard deviation in the fits, normalized to that of the optimal model, projected onto the E - M_{He} plane. We also show the constraints $M_{\text{ej}}/E_{\text{ej}} = \text{const}$ (blue) and $M_{\text{ej}}^2/E_{\text{ej}} = \text{const}$ (red) provided by the photospheric velocity evolution and the bolometric lightcurve, respectively.

5.4. Error sensitivity

Figure 15 shows the sensitivity of the derived quantities to errors in the distance, extinction and a systematic error in the photospheric velocity for our sample of Type IIb SNe. To compare with approximate scalings the derived quantities are expressed in terms of the mass and energy of the ejecta, but this does not affect the conclusions. For the helium-core mass and explosion energy the dependence on the distance and extinction is weak, whereas the dependence on the photospheric velocity is strong. For the mass of ^{56}Ni the sensitivity on the distance and extinc-

tion is strong, whereas the dependence on the photospheric velocity is weak. In general we see that an error in the distance and extinction mainly corresponds to an error in the mass of ^{56}Ni , whereas an error in the photospheric velocity mainly corresponds to an error in the helium-core mass and explosion energy. This behaviour is in agreement with the model grid dependencies discussed in Sect. 3.3 and with the qualitative discussion in E14a, based on the Arnett (1982) model. In Fig. 15 we show the scalings expected from the Arnett (1982) model and, as previ-

Table 3. Explosion energy, helium-core mass, mass and mixing of the ^{56}Ni and epoch of explosion for the best-fit models for the sample of individually studied Type IIb SNe.

SN	E (10^{51} erg)	M_{He} (M_{\odot})	M_{Ni} (M_{\odot})	Mix_{Ni}	JD_{exp} (+2400000) (days)
1996cb	1.45 (+1.25,-0.67)	5.12 (+1.38,-1.02)	0.106 (+0.175,-0.067)	1.00 (+0.00,-0.00)	50430.50 (+0.71,-1.12)
1993J	0.69 (+0.47,-0.33)	3.25 (+0.50,-0.57)	0.100 (+0.042,-0.023)	0.85 (+0.14,-0.07)	49074.00 (+0.00,-0.00)
2003bg	1.70 (+0.36,-0.80)	6.88 (+0.75,-1.53)	0.175 (+0.102,-0.071)	1.00 (+0.00,-0.00)	52689.50 (+0.00,-0.00)
2008ax	0.80 (+0.48,-0.35)	3.38 (+0.53,-0.55)	0.081 (+0.042,-0.041)	0.90 (+0.00,-0.11)	54528.80 (+0.00,-0.00)
2011dh	0.64 (+0.38,-0.30)	3.31 (+0.54,-0.57)	0.075 (+0.028,-0.020)	1.05 (+0.08,-0.00)	55713.00 (+0.00,-0.00)
2011ei	3.60 (+2.44,-1.25)	7.50 (+2.00,-0.76)	0.032 (+0.011,-0.011)	1.00 (+0.00,-0.00)	55763.00 (+1.12,-0.00)
2011fu	1.90 (+1.00,-0.87)	6.12 (+1.01,-1.15)	0.231 (+0.089,-0.101)	1.00 (+0.00,-0.00)	55820.50 (+1.12,-3.84)
2011hs	0.56 (+0.27,-0.27)	2.62 (+0.31,-0.31)	0.094 (+0.045,-0.040)	1.55 (+0.05,-0.78)	55874.50 (+2.55,-0.00)
2013df	1.11 (+1.08,-0.48)	3.69 (+1.20,-0.63)	0.056 (+0.006,-0.000)	0.80 (+0.60,-0.16)	56449.80 (+2.12,-0.00)

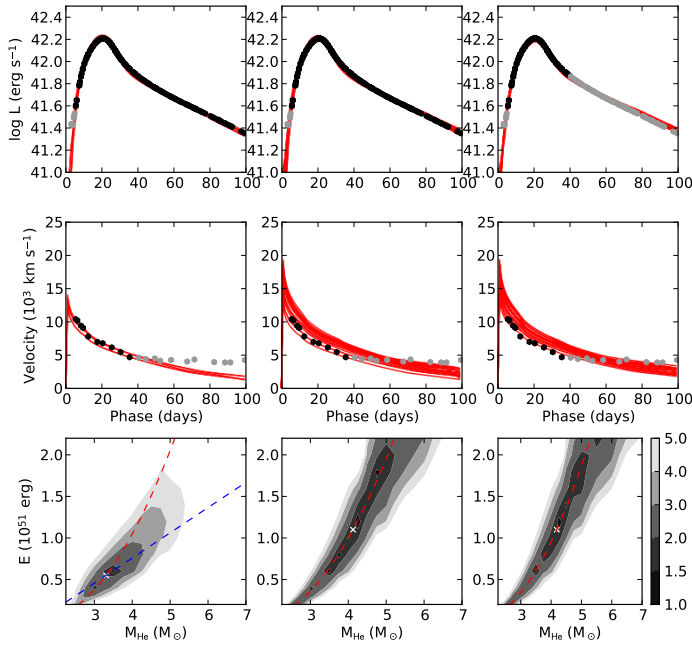


Fig. 16. Model bolometric lightcurve (upper panels) and photospheric velocity evolution (middle panels) as compared to the observed UV to MIR pseudo-bolometric lightcurve and estimated photospheric velocity evolution for SN 2011dh. Models with a normalized standard deviation <2 are shown in red for the cases when the lightcurve and photospheric velocity evolution (left panels), only the lightcurve (middle panels) and only the diffusion phase lightcurve (right panels) were used in the fit. The lower panels show the corresponding contour plots displayed as in Figs. 12 and 13.

ously discussed in Sect. 3.3, these are qualitatively followed by the model grid.

5.5. Degeneracy of the solution

Figure 16 shows the bolometric lightcurve and photospheric velocity evolution for models with a normalized standard deviation <2 as compared to observations for SN 2011dh. The left, middle and right panels show the cases when lightcurve and photospheric velocity evolution, only the lightcurve and only the dif-

fusion phase lightcurve were used in the fit, respectively. The lower panels show the corresponding contour plots and in the two latter cases, where the photospheric velocity evolution is not used in the fit, the solution is completely degenerate. This is not obvious as the diffusion phase and tail phase lightcurves might provide independent constraints arising from the diffusion time for thermal radiation and the optical depth for γ -rays, respectively. In E14a we argued that this could be the case, given that the diffusion time and optical depth provide the constraints $M_{\text{ej}}^3/E_{\text{ej}} = \text{const}$ and $M_{\text{ej}}^2/E_{\text{ej}} = \text{const}$, respectively, in the Arnett (1982) model. However, as discussed in Sect. 3.3, this model assumes a constant opacity and the constraint provided by diffusion phase lightcurve is rather $M_{\text{ej}}^2/E_{\text{ej}} = \text{const}$ for the hydrodynamical models. Therefore the diffusion phase and tail phase lightcurves appear to provide similar constraints and, as seen in the middle and right panels of Fig. 16, the degeneracy regions well follow the $M_{\text{ej}}^2/E_{\text{ej}} = \text{const}$ relation. The photospheric velocity evolution is expected to provide a constraint similar to $M_{\text{ej}}/E_{\text{ej}} = \text{const}$, which would break the degeneracy and, as seen in the left panels of Fig. 16, this is also the case. We have exemplified with SN 2011dh, but the conclusion is the same for the other SNe in the CSP and literature samples.

5.6. Sample statistics

Our sample of Type IIb SNe, in particular those that have been individually studied in the literature, is likely biased towards odd objects, and is neither time, nor volume-limited. On the other hand, the sample is the relatively large, and the modelling used relatively advanced. No parameter studies specifically aimed at Type IIb SNe have previously been published, but the stripped-envelope SN study by Lyman et al. (2014) includes 8 of the Type IIb SNe in our sample, as well as SN 2004el missing from our sample. Their results are based on the approximate (Arnett 1982) model, but are nevertheless interesting to compare with.

5.6.1. Parameter correlations

Figure 17 shows E versus M_{He} (left panel), M_{Ni} versus M_{He} (middle panel) and M_{Ni} versus E (right panel) for our sample of Type IIb SNe. The caveat that the explosion energy may be underestimated for SNe with relatively massive hydrogen envelopes (Sect. 4.1), and is also uncertain due to our simplified treatment of the opacity (Sect. 3.5), should be kept in mind.

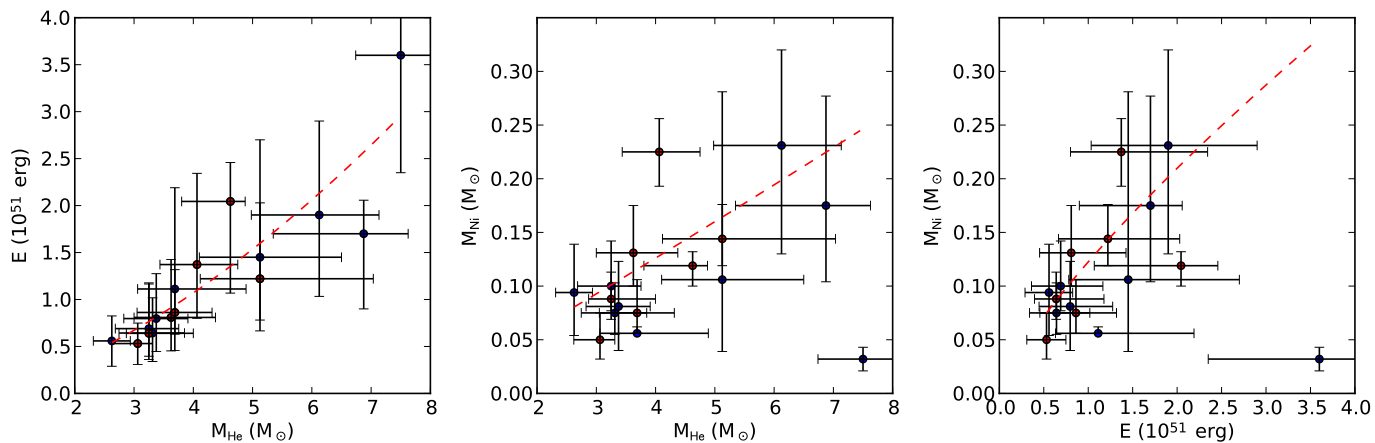


Fig. 17. E versus M_{He} (left panel), M_{Ni} versus M_{He} (middle panel) and M_{Ni} versus E (right panel) for our sample of Type IIB SNe (black circles), where we also show power-law fits as red dashed lines.

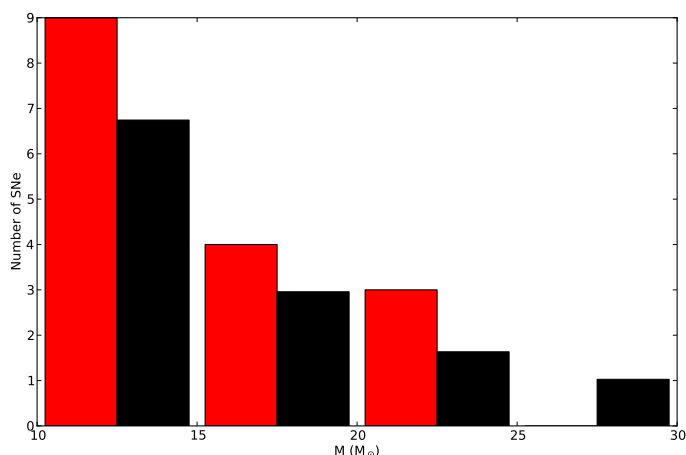


Fig. 18. Number of SNe with initial mass in the 8-15, 15-20, 20-25 and 25-30 M_{\odot} bins for our sample of Type IIB SNe (red) as compared to a standard Salpeter IMF (black).

We find a correlation between E and M_{He} , best fitted with a power law with index 1.8. Some caution is advised as this is similar to the $M_{\text{ej}}^2/E_{\text{ej}} = \text{const}$ degeneracy curve (Sect. 5.5). However, the SNe that obviously belongs to the low and high mass ends, like SNe 2011hs and 2003bg, well follows the relation. Lyman et al. (2014) also find a correlation, and assuming a mass for the compact remnant of 1.5 M_{\odot} and ignoring the difference between the explosion and ejecta energy, their results for the Type IIB SNe are best fitted with a power law with index of 1.8, in agreement with our results. The correlation found is also in qualitative agreement with the relation between progenitor mass and explosion energy for Type IIP SNe suggested by Poznanski (2013), which has a power law index of 3.

We find correlations between M_{Ni} and E and between M_{Ni} and M_{He} , best fitted with power laws with indices 0.7 and 1.1, respectively. Lyman et al. (2014) also find correlations between these quantities, and assuming a mass for the compact remnant of 1.5 M_{\odot} and ignoring the difference between the explosion and ejecta energy, their results for the Type IIB SNe are best fitted with power laws with indices of 0.89 and 1.36, respectively, in good agreement with our results. Hamuy (2003) finds a correlation between M_{Ni} and E for Type IIP SNe, using relations derived from model grid fits by Litvinova & Nadezhin (1983,

1985), and their results are best fitted with a power law with index 0.88, in good agreement with our results for Type IIB SNe. The correlations found are also in qualitative agreement with the relations between progenitor mass and M_{Ni} and expansion velocities and M_{Ni} for Type IIP SNe found by Fraser et al. (2011) and Maguire et al. (2012), respectively. An increase of the mass of ^{56}Ni with explosion energy is expected as this element is produced in the explosive nucleosynthesis, and such a trend is also predicted by modelling (e.g. Woosley & Weaver 1995). We note that SN 2011ei is an extreme outlier in the M_{Ni} versus M_{He} and M_{Ni} versus E plots, showing the highest helium-core mass and explosion energy, but the lowest mass of ^{56}Ni . This suggests a different nature of this SN, and a scenario where the inner parts of the ejecta fall back onto the compact remnant (e.g. Woosley & Weaver 1995), strongly reducing the amount of ejected ^{56}Ni , might be interesting to investigate.

5.6.2. Initial mass function

Figure 18 shows the number of SNe with initial mass in the 10-15, 15-20, 20-25 and 25-30 M_{\odot} bins for our sample of Type IIB SNe, as well as a standard Salpeter IMF with a minimum initial mass of 10 M_{\odot} . We find no SNe instead of the expected 4.6 SNe in the $>25 M_{\odot}$ bins, but except for this there are no significant deviations from a standard Salpeter IMF. We find 56 percent of the SNe in the $<15 M_{\odot}$ bin and 81 percent of the SNe in the $<20 M_{\odot}$ bins. The results for the Type IIB SNe from Lyman et al. (2014) are similar, and assuming a mass of the compact remnant of 1.5 M_{\odot} , their results corresponds to 67 percent of the SNe in the $<15 M_{\odot}$ bin and 100 percent of the SNe in the $<20 M_{\odot}$ bins. Although single-star mass-loss rates are uncertain (Smith 2014), recent stellar evolutionary models (e.g. Ekström et al. 2012), predicts the turning-point where the hydrogen envelope is lost to occur at $\sim 25 M_{\odot}$ at solar metallicity for single stars, in reasonable agreement with observations of galactic RSGs (e.g. Levesque et al. 2005) and WR stars (e.g. Hamann et al. 2006), respectively. Given this, the implication of our result is quite clear; either the binary channel is dominating the production of Type IIB SNe, or our understanding of single-star mass-loss is incomplete or even incorrect.

The idea that Type IIB SNe arise mainly from the binary channel has been proposed by several authors (references), but the size of the sample and the relatively detailed modelling used strengthen the support for this scenario considerably. Note that

a similar conclusion may hold for Type Ib and Ic SNe as well, and parameter studies using approximate lightcurve modelling (Cano 2013; Taddia et al. 2014; Lyman et al. 2014) find that the ejecta masses for these SNe are typically only a few solar masses and similar to those of Type IIb SNe. As shown by Lyman et al. (2014) using the stellar population synthesis code BPASS (Eldridge & Stanway 2009), this is considerably lower than expected for a population of single WR stars, and is also considerably lower than what would be expected for the observed population of galactic WR stars. A sample study of Type Ib and Ic SNe, using more detailed modelling, that could verify and possibly strengthen these results, would therefore be of great interest.

6. Conclusions

We present HYDE, a new 1-D hydrodynamical code, and use it to build a grid of SN models based on bare helium-core models evolved with MESA STAR. Such a grid is well suited to model the diffusion and early tail phase of Type IIb SNe, as the progenitors of these are thought to have lost all but a tiny fraction of their hydrogen envelopes. The dependences of the observed quantities on the progenitor and SN parameters are investigated, and found to be qualitatively similar to those of the approximate Arnett (1982) model. However, significant quantitative differences do exist, likely because of the constant opacity assumed in this model. Limitations in our method is discussed, in particular with respect to the opacity, where our simplified treatment results in a significant uncertainty in the estimated explosion energies, whereas the estimated helium-core masses are less affected. We also investigate the effects of a low-mass hydrogen envelope on the observed properties, and find these to be negligible after the luminosity minimum, expect for relatively massive hydrogen envelopes, where the photospheric velocities are decreased, likely because of deceleration of the helium core. This results in an significant uncertainty in the explosion energy for SNe with relatively massive hydrogen envelopes, and our method rather measures the explosion energy deposited in the helium core.

We use an automated fitting procedure to fit the bolometric lightcurves and photospheric velocities for our sample of Type IIb SNe to the grid of SN models. This allows us to take into account the uncertainties in distance, extinction and photospheric velocities, as well as to investigate the degeneracy of the solutions. The estimated progenitor and SN parameters for SNe 1993J, 2003bg, 2008ax, 2011dh and 2011hs are in reasonable agreement with hydrodynamical modelling in other works. However, in the case of 1993J and 2003bg, the derived explosion energies are significantly lower, likely due to the effect of their relatively massive hydrogen envelopes. We find an error in the distance and extinction to propagate mainly to derived the mass of ^{56}Ni , and a systematic error in the photospheric velocity to propagate mainly to the derived helium-core mass and explosion energy. If the photospheric velocities are not used in the fit, there is an almost complete degeneracy along the $M_{\text{ej}}^2/E = \text{const}$ curve, but when these are included the degeneracy is broken and the fit becomes quite robust.

We find correlations between the SN and progenitor parameters, the explosion energy increasing with helium-core mass, and the mass of ^{56}Ni increasing with the explosion energy. These correlations are best fitted with power-laws with indices of 1.8 and 1.1, respectively, in good agreement with the results in Lyman et al. (2014), obtained for a smaller sample using the approximate Arnett (1982) model. The initial masses of our sample of Type IIb SNe follows a standard Salpeter IMF reasonably well, although there is an under-population in the $>25 M_{\odot}$ range.

The fractions of SNe with initial masses $<15 M_{\odot}$ and $<20 M_{\odot}$ are 56 and 81 percent, respectively, in good agreement with the results in Lyman et al. (2014). Although single-star mass-loss rates are uncertain, single stars with initial masses much below $\sim 25 M_{\odot}$ are not expected to lose their hydrogen envelopes before core-collapse, and the implication of this result is quite clear; either the binary channel is dominating the production of Type IIb SNe, or our understanding of single-star mass-loss needs a serious revision. This conclusion is not new, and the evidence for this have been growing since the discovery of SN 1993J, but the size of the sample and the relatively detailed modelling used strengthen it considerably.

7. Acknowledgements

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Appendix A: The CSP sample of Type IIb SNe

The Type IIb SNe in the CSP sample consist of SNe 2004ex, 2004ff, 2005Q, 2006T, 2006ba, 2006bf, 2008aq, 2009K, 2009Z and 2009dq. The observational data for these SNe are described in Stritzinger (2015) and SNe 2006bf and 2009dq have been excluded due to bad sampling of the lightcurves. The photospheric velocities were estimated from the absorption minimum of the Fe II 5169 Å line, in turn measured from the observed spectra as described in Stritzinger (2015). Wherever used, the line-of-site extinction within the Milky way has been adopted from the Schlegel et al. (1998) extinction maps, recalibrated by Schlafly & Finkbeiner (2011).

SN 2004ex Discovered 2004 Oct 11.34 UT (Jacques et al. 2004) in NGC 182 at an apparent magnitude of 17.7, and the explosion epoch is constrained by non-detections from Oct 6.35 UT (<20.0 mag) and 10.33 (<19.0 mag) (Shimasaki et al. 2004). The photometric data covers the u to H bands and the 5-85 days period after discovery. Both the diffusion peak and early tail are well covered. The distance modulus for NGC 182 was adopted as the median and standard deviation of the literature values given by the NASA/IPAC Extragalactic Database (NED), being 34.83 ± 0.27 mag. The line-of-sight extinction within NGC 182 was estimated to $E(B-V)_{\text{H}} = 0.078$ mag by comparison of the $V-i$ colour at r -band maximum with SN 2011dh. Adding the line-of-sight-extinction within the Milky Way ($E(B-V)_{\text{MW}} = 0.021$ mag) gives $E(B-V) = 0.099$ mag. The pseudo-bolometric lightcurve was calculated using the $uBVriJH$ bands.

SN 2004ff Discovered 2004 Oct 30.40 UT (Pugh et al. 2004) in ESO 552-G40 at an apparent magnitude of 18.0, and the explosion epoch is constrained by non-detections from 2004 Oct 13.41 (<19.0 mag) and 21.40 UT (<18.0 mag) (Pugh et al. 2004). The photometric data covers the u to H band and the 5-80 days period after discovery. The rise to peak is not well covered and we have included observations from Pugh et al. (2004) to extend this coverage. In the absence of literature measurements of the distance to ESO 552-G40 we adopt the Virgo, Great Attractor and Shapley corrected kinematic distance modulus given by NED, being 34.82 ± 0.15 mag. The line-of-sight extinction

within ESO 552-G40 was estimated to $E(B-V)_H=0.138$ mag by comparison of the $V-i$ colour at r -band maximum with SN 2011dh. Adding the line-of-sight-extinction within the Milky Way ($E(B-V)_{MW}=0.029$ mag) gives $E(B-V)=0.167$ mag. The pseudo-bolometric lightcurve was calculated using the $uBVriJH$ bands.

SN 2005Q Discovered 2005 Jan 28.80 UT (Monard 2005) in ESO 244-G31 at an apparent magnitude of 17.2, and the explosion epoch is constrained by a non-detection from 2004 Dec 30.81 UT (<18.7 mag). The photometric data covers the u to i bands and the 0-50 days period after discovery. The rise to peak and the early tail is not well covered. The distance modulus for ESO 244-G31 was adopted as the median and standard deviation of the literature values given by NED, being 34.83 ± 0.18 mag. The $V-i$ colour at r -band maximum was bluer than for SN 2011dh, so the line-of-sight extinction within ESO 244-G31 was assumed to be negligible. The line-of-sight-extinction within the Milky Way is $E(B-V)_{MW}=0.023$ mag. The pseudo-bolometric lightcurve was calculated using the $uBVri$ bands.

SN 2006T Discovered 2006 Jan 30.99 UT (Monard 2006b) in NGC 3054 at an apparent magnitude of 17.2, and the explosion epoch is constrained by a non-detection from 2006 Jan 16.96 UT (<18.0 mag) (Monard 2006b). The photometric data covers the u to H bands and the 0-125 days period after discovery. Both the diffusion peak and the early tail are well covered. The distance modulus for NGC 3054 was adopted as the median and standard deviation of the literature values given by NED, being 32.58 ± 0.35 mag. The $V-i$ colour at r -band maximum was bluer than for SN 2011dh, so the line-of-sight extinction within NGC 3054 was assumed to be negligible. The line-of-sight-extinction within the Milky Way is $E(B-V)_{MW}=0.181$ mag. The pseudo-bolometric lightcurve was calculated using the $uBVriJH$ bands.

SN 2006ba Discovered 2006 Mar 19.81 UT (Monard 2006a) in NGC 2980 at an apparent magnitude of 17.7, and the explosion epoch is constrained by a non-detection from 2006 Feb 5.04 UT (<18.8) (Monard 2006a). The photometric data covers the u to H bands and the 5-80 days period after discovery. The rise to peak is not well covered. The distance modulus for NGC 2980 was adopted as the median and standard deviation of the literature values given by NED, being 34.48 ± 0.26 mag. The line-of-sight extinction within NGC 2980 was estimated to $E(B-V)_H=0.212$ by comparison of the $V-i$ colour at r -band maximum with SN 2011dh. Adding the line-of-sight-extinction within the Milky Way ($E(B-V)_{MW}=0.046$) gives $E(B-V)=0.258$. The pseudo-bolometric lightcurve was calculated using the $BVriJ$ bands.

SN 2008aq Discovered 2008 Feb 27.44 UT (Chu et al. 2008) in MCG -02-33-20 at an apparent magnitude of 16.3, and the explosion epoch is constrained by a non-detection from 2008 Feb 10.47 UT (<19.1 mag) (Chu et al. 2008). The photometric data covers the u to H bands and the 5-120 days period after discovery. The rise to peak is not well covered but we have included observations from Chu et al. (2008) and Brown et al. (2008) to extend this coverage. The distance modulus for MCG -02-33-20 was adopted as the median and standard deviation of the literature values given by NED, being $32.45\pm 0.43, 0.43$ mag. The $V-i$ colour at r -band maximum was bluer than for SN 2011dh

so the line-of-sight extinction within MCG -02-33-20 was assumed to be negligible. The line-of-sight-extinction within the Milky Way is $E(B-V)_{MW}=0.040$ mag. The pseudo-bolometric lightcurve was calculated using the $uBVriJH$ bands.

SN 2009K Discovered 2009 Jan 14.07 UT (Pignata et al. 2009) in NGC 1620 at an apparent magnitude of 14.9, and the explosion epoch is constrained by a non-detection from 2009 Jan 11.08 UT (<18.0 mag) (Pignata et al. 2009). The photometric data covers the B to H bands and the 0-50 days period after discovery. The early tail is not covered. The distance modulus for NGC 1620 was adopted as the median and standard deviation of the literature values given by NED, being 33.15 ± 0.22 mag. The line-of-sight extinction within NGC 1620 was estimated to $E(B-V)_H=0.057$ mag by comparison of the $V-i$ colour at r -band maximum with SN 2011dh. Adding the line-of-sight-extinction within the Milky Way ($E(B-V)_{MW}=0.051$ mag) gives $E(B-V)=0.108$ mag. The pseudo-bolometric lightcurve was calculated using the $uBVri$ bands.

SN 2009Z Discovered 2009 Feb 2.53 UT (Griffith et al. 2009) in SDSS J140153.80-012035.5 at an apparent magnitude of 18.1, and the explosion epoch is constrained by a non-detection from 2008 Jun-Jul (<19.4 mag) (Griffith et al. 2009). The photometric data covers u to i bands and the 5-85 days period after discovery, although additional NIR photometry was obtained at ~ 400 days. The rise to peak is not well covered but we have included observations from Griffith et al. (2009) to extend this coverage. In the absence of literature measurements of the distance to SDSS J140153.80-012035.5 we adopt the the Virgo, Great Attractor and Shapley corrected kinematic distance modulus given by NED, being 35.26 ± 0.15 mag. The $V-i$ colour at r -band maximum was bluer than for SN 2011dh so the line-of-sight extinction within SDSS J140153.80-012035.5 was assumed to be negligible. The line-of-sight-extinction within the Milky Way is $E(B-V)_{MW}=0.042$ mag. The pseudo-bolometric lightcurve was calculated using the $uBVri$ bands.

Appendix B: Individually studied Type IIb SNe

The sample of Type IIb SNe that have been individually studied in the literature consist of SNe 1993J (e.g. Richmond et al. 1994, 1996), 1996cb (Qiu et al. 1999), 2003bg (Hamuy et al. 2009), 2008ax (e.g. Taubenberger et al. 2011), 2011dh (e.g. E14a), 2011fu (Kumar et al. 2013), 2011hs (Bufano et al. 2014), 2011ei (Milisavljevic et al. 2013) and 2013df (Van Dyk et al. 2014), omitting here 2010as (Folatelli et al. 2014), for which the data was published after our work began. Wherever used, the line-of-site extinction within the Milky way has been adopted from the Schlegel et al. (1998) extinction maps, recalibrated by Schlafly & Finkbeiner (2011).

SN 1996cb Discovered 1996 Dec 15.71 UT (Nakano et al. 1996) in NGC 3510 at an apparent magnitude of 16.5, and the explosion epoch is constrained by a non-detection from 1996 Nov 29 UT (<19.0 mag) (Qiao et al. 1996). The photometric data was taken from Qiu et al. (1999) and covers the B to R bands and the 5-160 days period after discovery. To extend the rise to peak coverage we also included observations from (Nakano et al. 1996) and (Qiao et al. 1996). The distance modulus for NGC 3510 was adopted as the mean and standard deviation of the literature values given by NED being 30.57 ± 1.03 .

The total line-of-sight extinction was taken as the mean of the upper limit determined by Qiu et al. (1999) from comparison to SN 1993J ($E(B-V)=0.12$) and the line-of-sight extinction within the Milky Way ($E(B-V)=0.12$), giving $E(B-V)=0.073\pm0.047$. The pseudo-bolometric lightcurve was calculated using the *BVR* bands. Estimates of the photospheric velocities using SYNOW were adopted from Deng et al. (2001). Takáts & Vinkó (2012) showed that this method gives results similar to those obtained from the absorption minimum of Fe II 5169 Å line for a sample of Type IIP SNe.

SN 2003bg Discovered 2003 Feb 25.70 UT (Wood-Vasey et al. 2003) in MCG -05-10-15 at an apparent magnitude of 15.0, and the explosion epoch is constrained by a non-detection from 2002 Nov 7.0 UT (<18.0 mag) (Wood-Vasey et al. 2003). The photometric data was taken from Hamuy et al. (2009) and covers the *B* to *K* bands and the 5-325 days period after discovery. The distance modulus for the host galaxy MCG-05-10-015 was adopted from Kelson et al. (2000), and the total line-of-sight extinction taken as the mean of the line-of-sight extinction within the Milky Way ($E(B-V)=0.02$) and an assumed upper limit of 0.1 mag additional extinction, giving $E(B-V)=0.070\pm0.05$. The pseudo-bolometric lightcurve was calculated using the *BVRIZJK* bands. Estimates of the photospheric velocities using the Monte Carlo (MC) radiative transfer code by Mazzali & Lucy (1993); Lucy (1999); Mazzali (2000) was adopted from Mazzali et al. (2009).

SN 2011ei Discovered 2011 Jul 25.43 UT (Marples et al. 2011) in NGC 6925 at an apparent magnitude of 18.0, and the explosion epoch is constrained by a non-detection from 2011 Jun 23.58 UT (<19.1 mag) (Marples et al. 2011). The photometric data was taken from Milisavljevic et al. (2013) and covers *U* to *I* bands and the 0-50 days period. The distance modulus for the host galaxy NGC 6925 was adopted as the mean and standard deviation of the literature values given by NED, being 32.42 ± 0.27 . The total line-of-sight extinction was taken as the mean of the upper limit determined by Milisavljevic et al. (2013) ($E(B-V)=0.232$), who used the equivalent width of the interstellar Na I D interstellar absorption lines and the relation between this and $E(B-V)$ by Turatto et al. (2003), and the line-of-sight extinction within the Milky Way ($E(B-V)=0.052$), giving $E(B-V)=0.142\pm0.09$. The pseudo-bolometric lightcurve was calculated using the *UBVRI* bands. Estimates of the photospheric velocities using SYNOW was taken from Milisavljevic et al. (2013).

SN 2011fu Discovered 2011 Sep 21.04 UT (Ciabattari et al. 2011) in UGC 1626 at an apparent magnitude of 15.8, and the explosion epoch is constrained by a non-detection from 2011 Aug 10 UT (<18.8 mag) (Ciabattari et al. 2011). The photometric data covers the *U* to *I* bands and the 10-175 days period and was taken from Kumar et al. (2013). In the absence of literature measurements of the distance to the host galaxy UGC 1626 we adopt the Virgo, Great Attractor and Shapley corrected kinematic distance modulus given by NED, being 34.36, and assume an error in this estimate of 50 percent. The total line-of-sight extinction was adopted from Kumar et al. (2013), who used the equivalent width of the interstellar Na I D interstellar absorption lines and the relation between this and $E(B-V)$ by Munari & Zwitter (1997) to estimate $E(B-V)=0.22\pm0.11$. The pseudo-bolometric lightcurve was calculated using the *UBVRI* bands. Measure-

ments of the absorption minimum of the Fe II 5169 Å line was taken from Kumar et al. (2013).

SN 2011hs Discovered 2011 Nov 12.48 UT (Drescher et al. 2011) in IC 5267 at an apparent magnitude of 15.5, and the explosion epoch is constrained by a non-detection from 2011 Oct 9.574 UT (<18.7 mag) (Drescher et al. 2011). The photometric data covers the *U* to *K* bands and 0-120 days period and was taken from Bufano et al. (2014). The distance modulus for the host galaxy IC 5267 was adopted as the mean and standard deviation of the literature values given by NED, being 32.18 ± 0.41 . The total line-of-sight extinction was estimated to $E(B-V)_H=0.400$ by comparison of the *V-i* colour at *r*-band maximum with SN 2011dh. Spectroscopic data was taken from Bufano et al. (2014), and the photospheric velocities estimated from the absorption minimum of the Fe II 5169 Å line, in turn measured with a simple automated algorithm described in E14a.

SN 2013df Discovered 2013 Jun 7.87 UT (Ciabattari et al. 2013) in NGC 4414 at an apparent magnitude of $??$, and the explosion epoch is constrained by a non-detection from 2013 $??$ UT ($<?.=$ mag) (Ciabattari et al. 2013). The photometric data covers the *B* to *H* bands and the 5-65 days period and was taken from Van Dyk et al. (2014). The distance modulus for the host galaxy NGC 4414 was taken as the Cepheid based measurement by Freedman et al. (2001), being 31.10 ± 0.05 . The total line-of-sight extinction was adopted from Van Dyk et al. (2014), which used the equivalent width of the interstellar Na I D interstellar absorption lines and the relation between this and $E(B-V)$ by Poznanski et al. (2012) to estimate $E(B-V)=0.097\pm0.016$. The pseudo-bolometric lightcurve was calculated using the *BVRIZJH* bands. The spectroscopic data was taken from Van Dyk et al. (2014) and the photospheric velocities estimated from the absorption minimum of the Fe II 5169 Å line in turn measured with a simple automated algorithm described in E14a.

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